

THE ACTIVE CENTAURS

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ABSTRACT

The Centaurs are recent escapees from the Kuiper Belt that are destined either to meet fiery oblivion in the hot inner regions of the solar system or to be ejected to the interstellar medium by gravitational scattering from the giant planets. Dynamically evolved Centaurs, when captured by Jupiter and close enough to the Sun for near-surface water ice to sublimate, are conventionally labeled as “short-period” (specifically, Jupiter-family) comets. Remarkably, some Centaurs show comet-like activity even when far beyond the orbit of Jupiter, suggesting mass loss driven by a process other than the sublimation of water ice. We observed a sample of 23 Centaurs and found nine to be active, with mass-loss rates measured from several kg s^{-1} to several tonnes s^{-1} . Considered as a group, we find that the “active Centaurs” in our sample have perihelia smaller than the inactive Centaurs (median 5.9 AU versus 8.7 AU), and smaller than the median perihelion distance computed for all known Centaurs (12.4 AU). This suggests that their activity is thermally driven. We consider several possibilities for the origin of the mass loss from the active Centaurs. Most are too cold for activity at the observed levels to originate via the sublimation of crystalline water ice. Solid carbon monoxide and carbon dioxide have the opposite problem: they are so volatile that they should drive activity in Centaurs at much larger distances than observed. We consider the possibility that activity in the Centaurs is triggered by the conversion of amorphous ice into the crystalline form accompanied by the release of trapped gases, including carbon monoxide. By imposing the condition that crystallization should occur when the crystallization time is shorter than the orbital period we find a qualitative match to the perihelion distribution of the active Centaurs and conclude that the data are consistent with the hypothesis that the Centaurs contain amorphous ice.

Key words: comets: general – Kuiper Belt – minor planets, asteroids

Online-only material: color figures

1. INTRODUCTION

The Centaurs experience repeated close encounters with the giant planets and, as a result, follow orbits which are both chaotic and dynamically short-lived compared to the age of the solar system. Numerical simulations give half-lives from 0.5 Myr to 30 Myr, with average values increasing with semimajor axis and perihelion distance (Horner et al. 2004) and a population median lifetime near 10 Myr (Tiscareno & Malhotra 2003). The short dynamical lifetimes suggest that the Centaurs must be resupplied from a more stable source region elsewhere, if their population is to remain in steady state. This source region is most likely the Kuiper Belt, from which objects drift first into Neptune-crossing trajectories and then are scattered between the giant planets (see Morbidelli 2008 for a recent review of Kuiper Belt dynamics). The Kuiper Belt contains several dynamically distinct classes of body, and it is not clear precisely from where in the Belt the Centaurs are derived. The source is popularly identified with the so-called scattered Kuiper Belt objects but other trans-Neptunian populations, or even completely different source regions, may be implicated (see Volk & Malhotra 2008). The Centaurs suffer several fates. A few percent of the Centaurs impact the giant planets, or are broken apart by gravitational stresses exerted upon them by near-miss interactions with the giants (the famous comet D/Shoemaker-Levy 9 is the best example of a body suffering both of these fates). About two thirds are ultimately ejected from the solar system into interstellar space (Tiscareno & Malhotra 2003). The remaining one third become temporarily trapped in small, near-Sun orbits, where they sublimate and are conventionally labeled as “Jupiter family comets” (JFCs). The Centaurs, then, are scientifically interesting as a kind of intermediate population,

halfway between the distant, frigid, and relatively stable bodies in the Kuiper Belt and the rapidly sublimating, disintegrating comets of the hot inner regions (see Jewitt 2004 for a broad account of the interrelations between these objects).

Surprisingly, there is no uniformly accepted definition of the Centaurs. Some dynamicists classify all objects with perihelia between the orbits of Jupiter and Neptune as Centaurs, regardless of their semimajor axes. This definition is reasonable, given that the dynamics are dominated by perihelic interactions with the planets but some qualifications are needed, since many resonant Kuiper Belt objects (like 134340 Pluto), have perihelia inside Neptune’s orbit but are not usefully considered as Centaurs. A detailed classification system has been proposed by Horner et al. (2004) in which the Centaurs are labeled by the planets with which they most strongly interact. We here adopt a simpler operational definition, modified from Jewitt & Kalas (1998), and suggested by the prototype Centaur, (2060) Chiron (Kowal et al. 1979), whose dynamics are dominated by perihelic and aphelic interactions with giant planets. We define the Centaurs as bodies whose orbits meet the following conditions.

1. The perihelion distance, q , and the semimajor axis, a , satisfy $a_J < q < a_N$ and $a_J < a < a_N$, respectively, where $a_J = 5.2$ AU is the semimajor axis of Jupiter and $a_N = 30.0$ AU is the semimajor axis of Neptune.
2. The orbit is not in 1:1 mean-motion resonance with any planet.

This definition has the practical advantage of simplicity; it efficiently isolates a subset of the solar system small-body population whose members are short-lived by virtue of their gravitational interactions with the giant planets. At the time of writing there are, by the above definition, 92 known Centaurs.

Table 1
Journal of Observations

Object	UT Date	Tel	R_{au}^a	Δ_{au}^b	α (deg) ^c
C/2001 M10	2002 Sep 08	UH 2.2 m	6.056	5.349	4.16
C/2001 M10	2004 Oct 10	Keck 10 m	9.147	8.291	3.39
P/2004 A1(LONEOS)	2006 Aug 01	UH 2.2 m	6.079	5.961	9.62
39P/Oterma	2002 Sep 07	UH 2.2 m	5.484	4.684	6.95
29P/SW1	2002 Sep 08	UH 2.2 m	5.810	4.991	6.24
174P/Echeclus (60558)	2006 Feb 25	Keck 10 m	12.968	12.230	3.00
P/2005 T3 (Read)	2007 Sep 19	UH 2.2 m	6.442	6.368	8.97
P/2005 T3 (Read)	2007 Oct 12	Keck 10 m	6.459	6.026	8.27
P/2005 S2 (Skiff)	2007 Oct 12	Keck 10 m	6.552	5.866	6.70
165P/Linear	2003 Jan 08	Keck 10 m	8.300	8.123	6.75
(63252) 2001 BL41	2002 Feb 22	UH 2.2 m	8.456	7.547	2.79
(63252) 2001 BL41	2002 Feb 23	UH 2.2 m	8.458	7.555	2.89
2000 GM137	2001 Feb 16	UH 2.2 m	6.970	6.348	6.62
2000 GM137	2001 Feb 17	UH 2.2 m	6.970	6.336	6.54
2003 QD112	2004 Oct 10	Keck 10 m	11.625	10.692	1.82
(145486) 2005 UJ438	2007 Oct 12	Keck 10 m	9.280	8.609	4.73
(32532) Thereus	2007 Feb 20	Keck 10 m	10.792	11.233	4.60
166P/2001 T4	2001 Nov 18	UH 2.2 m	8.590	7.841	4.49
166P/2001 T4	2002 Sep 07	UH 2.2 m	8.573	7.937	5.43
(5145) Pholus	2004 Feb 17	Keck 10 m	18.538	18.488	3.05
(148975) 2001 XA255	2002 Dec 07	Keck 10 m	15.062	14.262	2.24
(148975) 2001 XA255	2003 Jan 09	Keck 10 m	14.965	13.982	0.14
167P/2004 PY42	2004 Oct 10	Keck 10 m	12.234	11.840	4.36
2006 SX368	2007 Sep 19	UH 2.2 m	12.398	11.703	3.44
(54598) Bienor	2006 Aug 23	UH 2.2 m	18.369	17.470	1.49
2002 PN34	2006 Aug 23	UH 2.2 m	14.661	13.678	0.95
2002 VG131	2002 Dec 07	Keck 10 m	14.870	14.431	3.24
2001 XZ255	2003 Jan 29	Keck 10 m	16.079	15.096	0.09
2001 XZ255	2004 Oct 10	Keck 10 m	16.167	16.385	3.43
(52975) Cyllarus	2004 Oct 10	Keck 10 m	21.089	20.542	2.30

Notes.

^a Heliocentric distance in AU.

^b Geocentric distance in AU.

^c Phase angle in degrees.

In this paper, we focus attention on Centaurs which display comet-like mass loss. Our purpose is to identify common properties of the activity and to seek evidence concerning its origin. The active Centaurs have, so far, escaped systematic study as a group. One likely reason is that the cometary classification afforded to such objects makes no distinction between them and other, long-studied comets in the Jupiter family, Halley family, and Long-Period populations. We seek answers to such questions as (1) How do the properties of active Centaurs compare with those of other comets and with Centaurs as a whole?, (2) Is mass loss implicated in the reportedly bimodal distribution of the optical colors of the Centaurs?, and (3) What drives the mass loss from the active Centaurs?

2. OBSERVATIONS

The imaging observations reported here were taken at the University of Hawaii 2.2 m telescope (UH 2.2 m) and the Keck 10 m telescope, both located atop Mauna Kea, Hawaii. At the UH 2.2 m we employed a Tektronix 2048 × 2048 pixel charge-coupled device (CCD) camera having an image scale of 0.219 arcsec (pixel)⁻¹ when mounted at the f/10 Cassegrain focus. At the Keck, we employed the Low Resolution Imaging Spectrograph (LRIS) which, in its modern form, employs independent red and blue channels separated by a dichroic beam-splitter (Oke et al. 1995). A 5000 Å dichroic was used. The red-side detector is a front-side illuminated Tektronix 2048 ×

Table 2
Orbital Parameters of the Observed Centaurs^a

Object	q (AU) ^b	a^c	e^d	i (deg) ^e	Perihelion
C/2001 M10	5.303	26.660	0.801	28.0	2001 Jun 21
P/2004 A1(LONEOS)	5.463	7.895	0.308	10.6	2004 Aug 25
39P/Oterma	5.471	7.256	0.246	1.9	2002 Dec 21
29P/SW1	5.722	5.986	0.044	9.4	2004 Jun 30
174P/Echeclus (60558)	5.849	10.740	0.455	4.3	2015 May 06
P/2005 T3 (Read)	6.202	7.507	0.174	6.3	2006 Jan 13
P/2005 S2 (Skiff)	6.398	7.964	0.197	3.1	2006 Jun 30
165P/Linear	6.830	18.03	0.621	15.9	2000 Jun 15
(63252) 2001 BL41	6.880	9.816	0.299	12.4	1998 Feb 09
2000 GM137	6.951	7.907	0.121	15.8	2002 Feb 24
2003 QD112	7.935	18.974	0.582	14.5	1998 May 29
145486	8.255	17.527	0.529	3.8	2010 Jul 01
(32532) Thereus	8.524	10.615	0.197	20.4	1999 Feb 15
166P/2001 T4	8.564	13.880	0.383	15.4	2002 May 20
(5145) Pholus	8.685	20.292	0.572	24.7	1991 Sep 24
(148975) 2001 XA255	9.387	30.183	0.689	12.7	2010 Jun 04
167P/2004 PY42	11.784	16.140	0.270	19.1	2001 Apr 10
2006 SX368	11.963	22.236	0.462	36.3	2010 May 16
(54598) Bienor	13.155	16.485	0.202	20.8	2027 Dec 19
2002 PN34	13.328	30.781	0.567	16.7	2001 Nov 11
2002 VG131	14.869	17.492	0.150	3.2	2002 Nov 09
2001 XZ255	15.353	15.910	0.035	2.6	1984 Jun 22
(52975) Cyllarus	16.286	26.438	0.384	12.6	1989 Oct 27

Notes.

^a Objects are listed in the order of increasing perihelion distance.

^b Perihelion distance in AU.

^c Semimajor axis in AU.

^d Orbital eccentricity.

^e Orbital inclination.

2048 pixel CCD camera, with image scale 0.210 arcsec (pixel)⁻¹ and two read-out amplifiers. The blue-side used two Marconi (E2V) CCDs, each 2048 × 4096 pixels in size with two readout amplifiers per chip, and with a scale of 0.135 arcsec (pixel)⁻¹. On the blue side we used a B filter with effective wavelength $\lambda_e = 4370$ Å and full width at half-maximum (FWHM) of $\Delta\lambda = 900$ Å. On the red side, the filters employed were V ($\lambda_e = 5473$ Å, $\Delta\lambda = 950$ Å), R ($\lambda_e = 6800$ Å, $\Delta\lambda = 1270$ Å), and I ($\lambda_e = 8330$ Å, $\Delta\lambda = 3130$ Å).

At both telescopes, we constructed flats by taking dithered images of the twilight sky and then median combining these images, after subtracting a bias frame and scaling them to a common value, to eliminate stars. For comparison, we also constructed “data flats” from the median combination of images of astronomical targets taken through the night, relying on the lack of spatial correlation between images to eliminate field stars. We found that the data-flats were superior to the twilight-sky flats in the B and I filters and we used them exclusively. In V and R the two types of flat were comparable in quality. Photometric calibration of the flattened images was obtained from measurements of Landolt standard stars (Landolt 1992). We used standard stars with optical colors similar to, or slightly redder than, the Sun, in order to minimize color terms. The faintest Landolt stars were used in order to minimize any possible shutter errors, and we experimented with different exposures to confirm that the shutter timing response was linear at the 1% (0.01 mag) level. Lastly, we avoided standard stars with listed photometric uncertainties greater than 0.01 mag.

We observed a total of 23 Centaurs. A journal of observations is given in Table 1 and the main orbital parameters of these Centaurs are listed in Table 2. Not all objects listed in Table 1

Table 3
R-band Photometry for Coma Search

Object	UT Date ^a	$m_{2.2}$ ^b	$m_{3.3}$ ^c	m_d ^d	$m_R(1, 1, 0)$ ^e	r_e (km) ^f
C/2001 M10	2002 Sep 08	18.85	18.53	20.02 ± 0.04	> 10.92	< 12
P/2004 A1	2006 Aug 01	19.76	19.44	20.95 ± 0.04	> 11.33	< 10
39P/Oterma	2002 Sep 07	21.83	21.75	24.64 ± 0.05	> 14.27	< 3
29P/SW1	2002 Sep 08	16.32	15.86	17.01 ± 0.05	> 8.52	< 35
174P/Echeclus (60558)	2006 Feb 25	20.01	19.69	21.19 ± 0.05	> 8.70	< 33
P/2005 S2 (Skiff)	2007 Oct 12	22.97	23.07	>26.18	14.54	2
(63252) 2001 BL41	2002 Feb 22	20.35	20.33	>23.85	11.03	11
(63252) 2001 BL41	2002 Feb 23	20.39	20.37	>25.17	11.06	11
2000 GM137	2001 Feb 16	22.61	22.63	>25.62	13.88	3
2000 GM137	2001 Feb 17	22.61	22.62	>27.42	13.89	3
2003 QD112	2004 Oct 10	22.32	22.01	23.55 ± 0.04	> 11.62	< 9
(145486) 2005 UJ438	2007 Oct 12	20.17	20.15	>23.87	10.25	16
(32532) Thereus	2007 Feb 20	19.87	19.87	>23.39	9.05	28
166P/2001 T4	2001 Nov 18	20.01	19.78	21.60 ± 0.05	> 10.47	< 14
166P/2001 T4	2002 Sep 07	19.65	19.27	20.63 ± 0.05	> 10.04	< 18
(5145) Pholus	2004 Feb 17	19.96	19.94	>25.51	7.00	72
(148975) 2001 XA255	2002 Dec 07	22.74	22.74	>26.93	10.82	12
(148975) 2001 XA255	2003 Jan 09	22.39	22.30	>26.20	10.74	13
167P/2004 PY42	2004 Oct 10	20.69	20.64	24.04 ± 0.05	> 9.50	< 23
2006 SX368	2007 Sep 19	20.15	20.08	23.8 ± 0.1	9.00	28
(54598) Bienor	2006 Aug 23	21.11	21.53	>25.90	8.38	38
2002 PN34	2006 Aug 23	19.92	19.89	>26.23	8.26	40
2002 VG131	2002 Dec 07	23.08	23.25	>26.13	11.10	11
2001 XZ255	2003 Jan 09	22.44	22.34	>26.14	10.48	14
2001 XZ255	2004 Oct 10	21.11	23.11	>26.26	8.66	34
(52975) Cyllarus	2004 Oct 10	21.67	21.66	>26.29	8.23	41

Notes.

^a UT Date of the observation.

^b Magnitude within a 2.2 arcsec radius aperture.

^c Magnitude within a 3.3 arcsec radius aperture.

^d Coma magnitude within the 2.2 to 3.3 arcsec annulus (Equation (1)).

^e Red magnitude of the nucleus, computed from $m_{2.2}$ corrected to unit heliocentric and geocentric distance and zero phase angle (Equation 2).

^f Effective radius of the nucleus in km assuming red geometric albedo $p_R = 0.1$. Limits are given where coma was detected.

appear in subsequent tables because of field-star contamination and other problems that caused us to reject some data from further consideration. The estimated semimajor axes of two of the objects in our sample (namely, (148975) 2001 XA255 with $a = 30.183$ AU and 2002 PN34 with $a = 30.781$) have been revised following the acquisition of observations, such that they (slightly) exceed the semimajor axis of Neptune (30.070 AU). We have elected to keep them in the sample, while noting that none of the conclusions drawn would be different if we instead chose to exclude them.

The next section gives brief introductions to the active Centaurs considered in this work.

2.1. 29P/Schwassmann–Wachmann 1

This object is labeled and traditionally regarded as a comet, but it also meets the dynamical definition of a Centaur and so it is appropriate to consider it here. Our optical observations (Table 3) show 29P in an active state, with a visible extent much larger than 100 arcsec (Figure 1). The colors are slightly redder than the Sun and essentially constant with respect to aperture size, except for a marginally significant trend towards bluer colors at larger radii. Earlier observers have reported that the broadband colors of 29P measured within a very large (88 arcsec diameter) aperture vary with the level of activity, becoming redder as the comet brightens (Kiselev & Chernova 1979).

Since discovery in 1927, 29P has shown repeated photometric outbursts superimposed upon a platform of constant activity

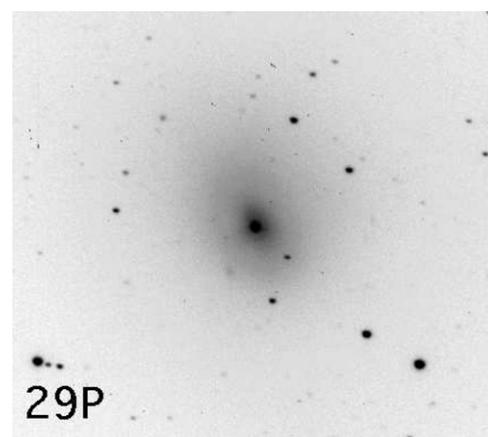


Figure 1. 29P/Schwassmann–Wachmann 1 imaged in the R-band filter on UT 2002 September 8. This is a 300 s integration with the University of Hawaii 2.2 m telescope. The telescope was tracked nonsidereally to follow the motion of the Centaur. Field shown has north to the bottom, east to the right, and is 130 arcsec wide. The centrally located Centaur is marked with a bar.

(e.g., Roemer 1958; Jewitt 1990; Trigo-Rodríguez et al. 2008). Spectroscopic measurements show that the activity is driven by the sublimation of carbon monoxide (CO) ice (Senay & Jewitt 1994; Crovisier et al. 1995), with CO mass-loss rates of 1000 to 2000 kg s⁻¹ (c.f., Table 4). Other species have been reported when the comet is at its brightest, including CO⁺ and CN (Cochran & Cochran 1991).

Table 4
Model Mass-Loss Rates

Object	R (AU) ^a	Δ (AU) ^b	α (deg) ^c	m_d ^d	C_d (m ²) ^e	τ (s) ^f	M_d (kg) ^g	$\frac{dM_d}{dt}$ ^h
C/2001 M10	6.056	5.349	4.16	20.0	5.77×10^7	8.16×10^3	2.31×10^6	2.8×10^2
P/2004 A1	6.079	5.961	9.62	21.0	4.55×10^7	9.10×10^3	1.82×10^6	2.0×10^2
39P/Oterma	5.484	4.684	6.95	24.6	7.61×10^5	6.97×10^3	3.04×10^4	4.3×10^0
29P/SW1	5.810	4.991	6.24	17.0	9.61×10^8	7.53×10^3	3.84×10^7	5.1×10^3
174P	12.970	12.230	3.00	21.2	1.85×10^8	2.26×10^4	7.41×10^6	3.2×10^2
P/2005 S2	6.552	5.866	6.70	> 26.28	$< 2.93 \times 10^5$	9.12×10^3	$< 1.18 \times 10^4$	$< 1.3 \times 10^0$
63252	8.458	7.555	2.89	> 25.17	$< 1.23 \times 10^6$	1.25×10^4	$< 4.93 \times 10^4$	$< 3.9 \times 10^0$
2000 GM137	6.970	6.336	6.54	> 27.42	$< 1.24 \times 10^5$	1.00×10^4	$< 4.95 \times 10^3$	$< 0.5 \times 10^0$
2003 QD112	11.625	10.692	1.82	23.5	1.33×10^7	1.92×10^4	5.31×10^5	2.8×10^1
145486	9.280	8.609	4.73	> 23.87	$< 6.63 \times 10^6$	1.46×10^4	$< 2.65 \times 10^5$	$< 1.8 \times 10^1$
32532	10.792	11.233	4.60	> 23.39	$< 1.80 \times 10^7$	1.98×10^4	$< 7.19 \times 10^5$	$< 3.6 \times 10^1$
166P/2001 T4	8.573	7.937	5.43	20.6	1.10×10^8	1.32×10^4	4.40×10^6	3.3×10^2
(5145) Pholus	18.538	18.488	3.05	> 25.51	$< 1.08 \times 10^7$	3.73×10^4	$< 4.30 \times 10^5$	$< 1.2 \times 10^1$
148975	15.062	14.262	2.24	> 26.93	$< 1.37 \times 10^6$	2.73×10^4	$< 5.49 \times 10^4$	$< 2.0 \times 10^0$
167P/2004 PY42	12.234	11.840	4.36	24.0	1.31×10^7	2.15×10^4	5.24×10^5	2.4×10^1
2006 SX368	12.398	11.703	3.44	23.8	1.52×10^7	2.13×10^4	6.10×10^5	2.9×10^1
54598	18.369	17.470	1.49	> 25.90	$< 6.03 \times 10^6$	3.52×10^4	$< 2.41 \times 10^5$	$< 6.9 \times 10^0$
2002 PN34	14.661	13.678	0.95	> 26.23	$< 2.11 \times 10^6$	2.60×10^4	$< 8.45 \times 10^4$	$< 3.2 \times 10^0$
2002 VG131	14.870	14.431	3.24	> 26.13	$< 3.10 \times 10^6$	2.75×10^4	$< 1.24 \times 10^5$	$< 4.5 \times 10^0$
2001 XZ255	16.167	16.385	3.43	> 26.26	$< 3.76 \times 10^6$	3.19×10^4	$< 1.50 \times 10^5$	$< 4.7 \times 10^0$
52975	21.089	20.542	2.30	> 26.29	$< 7.03 \times 10^6$	4.28×10^4	$< 2.81 \times 10^5$	$< 6.6 \times 10^0$

Notes.

^a Heliocentric distance.

^b Geocentric distance.

^c Phase angle.

^d Coma annulus magnitude from Table 3. We have taken the most sensitive limit for each object in which a coma was not detected.

^e Derived dust cross-section (Equation 3).

^f Photometry annulus crossing time (Equation 6).

^g Derived coma mass (kg), from C_d and Equation A4.

^h Derived mass-loss rate (kg s^{-1}), from Columns (g) and (f).

Thermal measurements give an upper limit to the nucleus radius of 27 ± 5 km and a lower limit to the geometric albedo of $p_R = 0.025 \pm 0.01$ (Stansberry et al. 2004). (These values are limits because of the substantial contamination of the infrared data by coma, and the need to make a model-dependent correction for coma.)

2.2. 39P/Oterma

The orbital parameters of this Centaur evolve substantially on timescales of decades owing to close interactions with Jupiter. For example, the perihelion was raised from 3.39 AU to 5.47 AU by a close approach (0.1 AU) to Jupiter in 1963, after which 39P was not re-observed until 1998. In our data from the UH 2.2 m telescope taken 2002 September 7, 39P was found to show a weak coma, barely visible in Figure 2. The object is slightly redder than the Sun, with no compelling evidence for a color gradient within the apertures employed for photometry (Table 5).

2.3. 165P/Linear

Coma was reported around this object (formerly known as 2000 B4) in observations taken with a 0.65 m diameter telescope on UT 2000 February 10 (Kusnirak & Balam 2000) but no coma has been reported since this initial report. Our own attempts to observe the object were made on UT 2002 January 9 using the UH 2.2 m telescope and on UT 2003 January 8, using the Keck 10 m telescope. The object was not detected down to approximate limiting magnitudes $m_R = 23$ in the former case and 24.5 in the latter. We conclude that either the object was

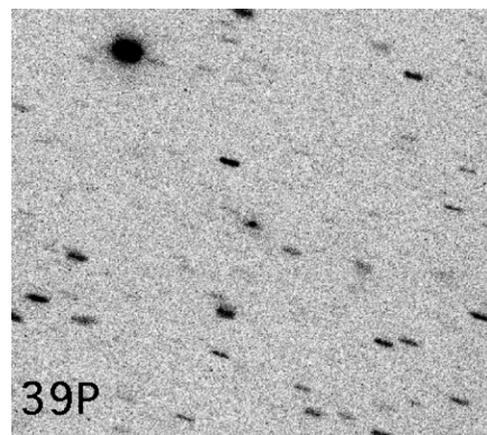


Figure 2. 39P/Oterma imaged in the R -band filter on UT 2002 September 8. This is the sum of two consecutive 900 s integrations with the University of Hawaii 2.2 m telescope. The telescope was tracked nonideally to follow the motion of the Centaur. Field shown has north to the bottom, east to the right, and is 130 arcsec wide.

very faint or the ephemeris position was in error by more than the field radius (about 3.5 arcminutes).

2.4. 166P/2001 T4

Our imaging data show a clear coma at all epochs (Figure 3) and the $BVRI$ color data show that 166P is optically very red. For example, in 2001 November (Table 5) we find $B - V = 0.87$, $V - R = 0.69$, $R - I = 0.78$ within a 5 pixel (1.1 arcsec) radius aperture. Bauer et al. (2003) reported 166P to be one of

Table 5
Color and Multi-Aperture Photometry

Object	UT 2005	ϕ^a	R^b	$B - V^c$	$V - R^c$	$R - I^c$
29P/SW1	2002 Sep 08	1.1	16.56 ± 0.01	0.79 ± 0.03	0.51 ± 0.03	0.52 ± 0.03
29P/SW1	2002 Sep 08	2.2	16.32 ± 0.01	0.78 ± 0.03	0.50 ± 0.03	0.52 ± 0.03
29P/SW1	2002 Sep 08	3.3	15.86 ± 0.01	0.79 ± 0.04	0.47 ± 0.03	0.50 ± 0.03
29P/SW1	2002 Sep 08	6.6	15.83 ± 0.01	0.75 ± 0.03	0.51 ± 0.03	0.48 ± 0.03
39P/Oterma	2002 Sep 07	1.1	22.08 ± 0.05	0.89 ± 0.07	0.41 ± 0.05	0.47 ± 0.10
39P/Oterma	2002 Sep 07	2.2	21.83 ± 0.05	0.74 ± 0.07	0.45 ± 0.05	0.59 ± 0.10
39P/Oterma	2002 Sep 07	3.3	21.75 ± 0.05	0.72 ± 0.07	0.40 ± 0.05	0.44 ± 0.10
166P/2001 T4	2001 Nov 18	1.1	20.47 ± 0.03	0.87 ± 0.04	0.69 ± 0.03	0.78 ± 0.04
166P/2001 T4	2001 Nov 18	2.2	20.01 ± 0.03	0.87 ± 0.04	0.73 ± 0.03	0.71 ± 0.04
166P/2001 T4	2001 Nov 18	3.3	19.78 ± 0.03	0.87 ± 0.04	0.66 ± 0.04	0.73 ± 0.04
166P/2001 T4	2001 Nov 18	6.6	19.35 ± 0.02	0.70 ± 0.04	0.73 ± 0.03	...
166P/2001 T4	2002 Sep 07	1.1	19.91 ± 0.03	...	0.70 ± 0.03	0.69 ± 0.04
166P/2001 T4	2002 Sep 07	2.2	19.65 ± 0.03	...	0.63 ± 0.03	0.70 ± 0.04
166P/2001 T4	2002 Sep 07	3.3	19.27 ± 0.03	...	0.52 ± 0.03	0.76 ± 0.04
166P/2001 T4	2002 Sep 07	3.3	19.03 ± 0.02	...	0.52 ± 0.03	0.76 ± 0.04
167P/2004 PY42	2004 Oct 10	1.1	20.80 ± 0.03	0.77 ± 0.03	0.55 ± 0.03	0.54 ± 0.02
167P/2004 PY42	2004 Oct 10	2.2	20.69 ± 0.02	0.80 ± 0.03	0.51 ± 0.03	0.51 ± 0.03
167P/2004 PY42	2004 Oct 10	3.3	20.64 ± 0.04	0.76 ± 0.03	0.49 ± 0.04	0.51 ± 0.04
C/2001 M10	2002 Sep 08	1.1	19.14 ± 0.02	1.22 ± 0.08 ^d
C/2001 M10	2002 Sep 08	2.2	18.85 ± 0.02	1.19 ± 0.08 ^d
C/2001 M10	2002 Sep 08	3.3	18.53 ± 0.03	1.17 ± 0.08 ^d
C/2001 M10	2002 Sep 08	6.6	18.52 ± 0.04	1.16 ± 0.08 ^d
C/2001 M10	2004 Oct 10	1.1	22.76 ± 0.06	0.77 ± 0.06	0.35 ± 0.06	0.52 ± 0.08
C/2001 M10	2004 Oct 10	2.2	22.08 ± 0.06	0.66 ± 0.10	0.49 ± 0.10	0.42 ± 0.08
174P/Echeclus (60558)	2006 02 25	1.1	20.27 ± 0.02
174P/Echeclus (60558)	2006 02 25	2.2	20.01 ± 0.02
174P/Echeclus (60558)	2006 02 25	3.3	19.69 ± 0.02
174P/Echeclus (60558)	2006 02 25	6.6	18.87 ± 0.02
2003 QD112	2004 10 10	1.1	22.49 ± 0.03
2003 QD112	2004 10 10	2.2	22.32 ± 0.03
2003 QD112	2004 10 10	3.3	22.18 ± 0.03
2003 QD112	2004 10 10	6.6	22.01 ± 0.05
P/2004 A1	2006 08 01	1.1	20.35 ± 0.02
P/2004 A1	2006 08 01	2.2	19.76 ± 0.02
P/2004 A1	2006 08 01	3.3	19.44 ± 0.02	...	0.32 ± 0.03	...
P/2004 A1	2006 08 01	6.6	19.01 ± 0.02	...	0.39 ± 0.03	...

Notes.

^a Radius of the photometry aperture, in arcseconds.

^b Apparent R -band magnitude and its 1σ uncertainty.

^c Color indices within each aperture.

^d The $B - R$ color is reported on this object only.

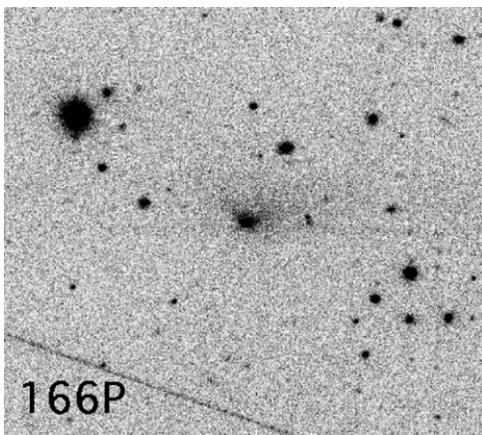


Figure 3. 166P/2001 T4 imaged in the R -band filter on UT 2002 September 7. This is a 300 s integration with the University of Hawaii 2.2 m telescope. The telescope was tracked nonsidereally to follow the motion of the Centaur. The diagonal line is from a passing satellite. Field shown has north to the bottom, east to the right, and is 130 arcsec wide.

the reddest objects in the solar system, with $V - R = 0.95 \pm 0.04$ and $R - I = 0.69 \pm 0.07$ measured in 2001 October and 2002 January within a 1.5 arcsec radius aperture. The reason for the difference is not clear: we used the same telescope and filters and our measurements were obtained within a month of Bauer et al.'s first observations. These authors also reported a radial color gradient in 166P, with the image core being substantially redder than the coma. Our data confirm the existence of such a radial color gradient (Table 5).

2.5. 167P/2004 PY42

This Centaur was observed on UT 2004 October 10 at the Keck 10 m telescope under photometric conditions. A weak coma was apparent in the raw data, and can be seen in Figure 4 (note that the diffuse object to the north east of 167P in the figure is a background galaxy). The coma is faint compared to the nucleus.

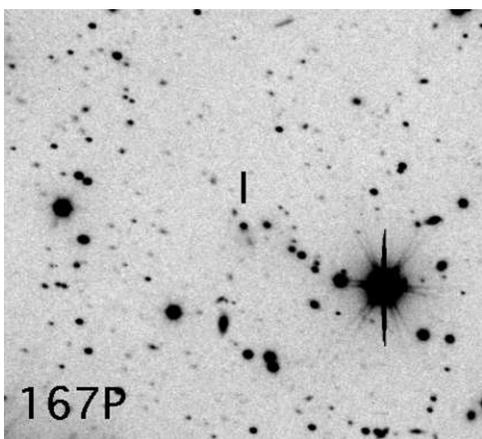


Figure 4. 167P/2004 PY42 imaged in the *R*-band filter on UT 2002 September 8. This is the sum of two 200 s integrations with the Keck 10 m telescope. The telescope was tracked considerably to follow the motion of the Centaur. Field shown has north to the bottom, east to the right, and is 130 arcsec wide.

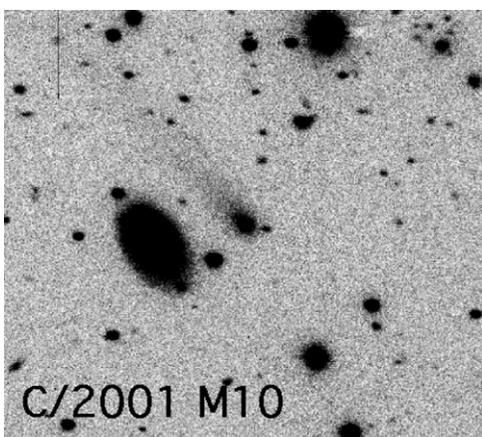


Figure 5. C/2001 M10 imaged in the *R*-band filter on UT 2002 September 8. This is a 300 s integration with the University of Hawaii 2.2 m telescope. The telescope was tracked considerably to follow the motion of the Centaur. Field shown has north to the bottom, east to the right, and is 130 arcsec wide.

2.6. C/2001 M10 (*Neat*)

This Centaur was observed on UT 2002 September 8 at the UH 2.2 m telescope and on UT 2004 October 10 using the Keck 10 m telescope (Figure 5). As a result of instrumental problems, observations on the former date were obtained only in the *B* and *R* filters, while on the latter date full *BVRI* colors were secured. The object faded by ~ 3.5 magnitudes between 2002 and 2004, as the heliocentric distance increased from 6.056 AU to 9.145 AU. Only ~ 2 magnitudes of this fading can be attributed to geometric effects, the remaining ~ 1.5 mag reflects a real decrease in the coma cross-section of C/2001 M10. However, the *B* – *R* color indices at the two epochs of observation, 1.22 ± 0.08 and 1.12 ± 0.09 , are completely statistically consistent within the measurement uncertainties. Since the coma contribution to the signal in the UT 2004 October 10 data was reduced relative to that from the nucleus by a factor ~ 4 , this suggests either that the nucleus contributes a negligible fraction of the total cross-section even at 9 AU or that the nucleus and the coma have the same *B* – *R* color.

2.7. P/2004 A1 (*Loneos*)

This object passed within 0.03 AU of Saturn in 1992 causing a sudden decrease in the perihelion distance from ~ 10 AU

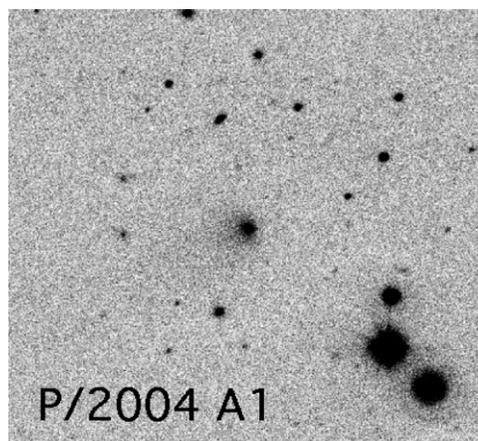


Figure 6. P/2004 A1 (*Loneos*) imaged in the *R*-band filter on UT 2006 July 1. This is a 300 s integration with the University of Hawaii 2.2 m telescope. The telescope was tracked considerably to follow the motion of the Centaur. Field shown has north to the bottom, east to the right, and is 130 arcsec wide.

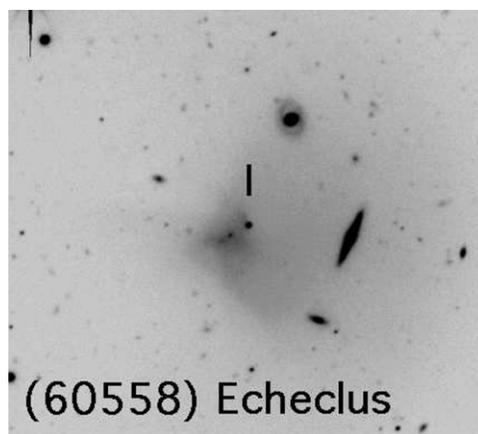


Figure 7. 174P/Echeclus (60558) imaged in the *R*-band filter on UT 2005 February 25. This is a 300 s integration with the Keck 10 m telescope. The telescope was tracked considerably to follow the motion of the Centaur. Field shown has north to the bottom, east to the right, and is 130 arcsec wide. The primary nucleus is marked with a bar.

to ~ 5 AU (Hahn et al. 2006). We observed it post-perihelion on UT 2006 July 01 at $R = 6.079$ AU, when it displayed a prominent coma (Figure 6). The absolute magnitude we derived, $m_R(1, 1, 0) > 10.92$, is consistent with the faintest values reported by Mazzotta et al. (2006).

2.8. 174P/Echeclus (60558)

The object (formerly asteroid 2000 EC98) was discovered and measured as an inert Centaur (Rousselot et al. 2005) but promptly erupted into activity in late 2005 (Choi et al. 2006). We imaged the object using Keck on 2006 February 25 (Figure 7). The coma showed a complicated post-outburst structure, visible up to 65 arcsec from the nucleus. We digitally removed background stars from the images and determined magnitudes within the standard apertures. The magnitude of the nucleus corrected to unit heliocentric and geocentric distances and to zero degrees phase angle, $m_R(1, 1, 0) > 8.7$ (Table 3), is consistent with the pre-activity value, $H_R = 9.03 \pm 0.01$, measured by Rousselot et al. Only *R* filter color measurements were possible at Keck due to time constraints (Table 5). However, spatially resolved colors were determined on UT 2006 February 24 by Bauer et al. (2008). They found a coma slightly

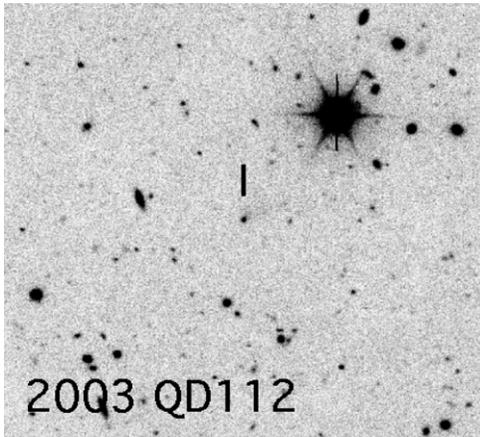


Figure 8. 2003 QD112 imaged in the *R*-band filter on UT 2004 October 10. This is a 300 s integration with the Keck 10 m telescope. The telescope was tracked considerably to follow the motion of the Centaur. Field shown has north to the bottom, east to the right, and is 130 arcsec wide. The centrally located Centaur is marked with a bar.

redder than sunlight with no substantial color gradients in the 5 arcsec to 30 arcsec projected radius range.

The nature of this object remains a mystery. The motion of the active secondary component suggests that it is a fragment ejected from the parent body, but neither radiation forces nor nongravitational forces due to asymmetric outgassing seem capable of producing the large relative velocity of separation (Weissman et al. 2006). Thermal infrared measurements show that the coma particles are large and emitted continuously, not impulsively (Bauer et al. 2008).

2.9. (5145) Pholus

(5145) Pholus (formerly 1992 AD) is well known for its ultra-red optical reflection spectrum and presumed organic composition. No coma has ever been reported around this object and no coma was detected in our Keck images taken UT 2004 February 17 (Table 3).

2.10. (63252) 2001 BL41

This object was observed as a point source on UT 2002 February 22 and 23 at the UH 2.2 m telescope with very similar apparent magnitudes on the two nights (Table 3).

2.11. 2003 QD112

We observed 2003 QD112 at the Keck telescope on UT 2004 October 10 when at 12.65 AU heliocentric distance and discovered a coma about this object (Figure 8). Only *R*-band data were secured (Table 3). We reimaged 2003 QD112 at the Keck on UT 2007 October 12. A coma was again evident in a blinked pair of 300 s *R*-band integrations, but the comet appeared in projection near to a much brighter field galaxy, making accurate photometry impossible on this date.

2.12. P/2005 S2 (Skiff)

We attempted to image P/2005 S2 on UT 2007 September 19 using the UH 2.2 m telescope. The object was not detected in 300 s *R*-band exposures with a limiting magnitude near $R = 23.5$, leading us to conclude either that the brightness was less than expected or that the ephemeris was substantially in error. Subsequently, we detected P/2005 S2 as a point-source object near magnitude 23.0 (Table 3) using the Keck 10 m on UT 2007 October 12.

2.13. P/2005 T3 (Read)

This Centaur was not detected on either of two attempts (see Table 1) with two different telescopes. On UT 2007 September 17 the object was projected against a background H II nebula which may have impeded its detection. On UT 2007 October 12 a bright background star filled the field of view with scattered light.

3. OBSERVATIONAL RESULTS

3.1. Coma Search

We first visually searched each Centaur for coma by eye, comparing images taken at different times in order to distinguish faint near-nucleus structure from occasional background objects. We also used multi-aperture photometry to measure, or set limits to, light from near-nucleus coma. Given a point-source nucleus embedded in an extended coma, the inner aperture of radius ϕ_1 would sample the Centaur nucleus combined with an uncertain contribution from near-nucleus coma while the annulus between ϕ_1 and an outer aperture of radius ϕ_2 would provide a measure of pure coma. Then, the apparent magnitude of the coma in the annulus between ϕ_1 and ϕ_2 is given by

$$m_d = -2.5 \log_{10} (10^{-0.4m_2} - 10^{-0.4m_1}) \quad (1)$$

in which m_1 and m_2 are the magnitudes within apertures of radii ϕ_1 and ϕ_2 . In practice, a small correction is needed because the point-spread function (PSF) of the image is nonzero in the annulus between ϕ_1 and ϕ_2 . We measured the magnitude of this correction to the brightness in the coma annulus using the PSFs of field stars in the same images as the Centaurs.

Based on trial and error, we selected aperture radii $\phi_1 = 2.2$ arcsec and $\phi_2 = 3.3$ arcsec as standard for our coma-search measurements. Smaller values of ϕ_1 proved overly sensitive to the steep surface brightness gradients found near the image core and to small differences in the PSF caused by relative motion between the Centaurs and the field stars used for reference. Conversely, substantially larger apertures, ϕ_2 , were more frequently contaminated by faint background sources and more subject to photometric errors from the uncertainty in the brightness of the night sky. With this choice of ϕ_1 and ϕ_2 , we found that the wings of the PSF between $\phi_1 \leq \phi \leq \phi_2$ typically contained only $\sim 1\%$ to 2% of the light with $\phi < \phi_1$.

The resulting coma annulus magnitudes, m_d , are listed in Column 5 of Table 3. Where no significant coma is detected, we have listed a 3σ limit to m_d computed from the scatter in multi-annulus measurements of ~ 5 to ~ 10 field stars. Coma is detected in nine of the 23 Centaurs observed. In eight of these nine cases, coma that was detected through multi-aperture photometry and Equation (1) was also visible to the eye in highly stretched images blinked rapidly on a computer. Only in Centaur 2006 SX368 did we find evidence for coma from the photometry that could not be corroborated visually.

3.2. Absolute Magnitudes and Sizes

The magnitudes of the Centaurs are corrected to $R_{\text{au}} = \Delta_{\text{au}} = 1$ and phase angle $\alpha = 0$ using

$$m_R(1, 1, 0) = m_R - 5 \log_{10}(R_{\text{au}}\Delta_{\text{au}}) - \Phi(\alpha) \quad (2)$$

where R_{au} and Δ_{au} are the heliocentric and geocentric distances in AU and $\Phi(\alpha)$ is the so-called phase function correction at phase angle α . The phase angles, α , at which the Centaurs were

observed span the range from $0 \leq \alpha \leq 10^\circ$ (Table 1). To first order, the phase function over this modest range of angles can be represented as having a linear (in magnitude space) dependence on α , given by $\Phi(\alpha) = 10^{(-0.4\beta\alpha)}$, where β [mag deg $^{-1}$] is a constant. However, at small phase angles, the phase functions of many objects display a narrow-angle peak (the “opposition-surge”) while at larger phase angles the phase function often appears measurably nonlinear. The phase functions of the active Centaurs are further uncertain, because the angle dependence of the scattering from the dust particles in the coma is not necessarily like that of the solid surface of the nucleus from which those particles were ejected. Here, we used the “HG” formalism of Bowell et al. (1989) with scattering parameter $G = 0.15$. This choice gives phase darkening similar to the constant $\beta = 0.1$ mag deg $^{-1}$ case (e.g., the Bowell formalism gives phase darkening at $\alpha = 10^\circ$ to 35% of the opposition value while $\beta = 0.1$ mag deg $^{-1}$ gives a 40% darkening).

The absolute magnitudes, $m_R(1, 1, 0)$, are listed in Table 3. The table also lists effective object radii, r_e (km), computed under the assumption of geometric albedo $p_R = 0.1$. The statistical uncertainty on $m_R(1, 1, 0)$ is the same as on m_R but the actual uncertainty is larger, owing to uncertainty introduced by the phase angle correction. The uncertainty on r_e is potentially large and unknown, pending future determinations of the albedos and phase functions of the observed objects. We include $m_R(1, 1, 0)$ and r_e here mainly to provide an approximate indication of the relative sizes of the observed objects. Note that entries for active Centaurs in Table 3 are presented as lower and upper limits on $m_R(1, 1, 0)$ and r_e , respectively, with no attempt to remove coma from the central aperture using profile-fitting or other techniques. The smallest observed Centaurs, P/2005 S2 ($r_e = 2$ km), 2000 GM137 ($r_e = 3$ km) and 39P ($r_e < 3$ km) have sizes consistent with the nuclei of better-observed JFCs. However, the median radius of the observed (inactive) Centaurs is $r_e = 15$ km (Table 3), about an order of magnitude larger than the characteristic radii of well-measured Jupiter family comet nuclei (Lamy et al. 2004) and an order of magnitude smaller than the characteristic radii of Kuiper Belt objects for which reliable physical data exist. These differences simply reflect the relative mean distances of the different types of object and the limitations of magnitude-limited photometry.

3.3. Colors

Like the Kuiper Belt objects, the Centaurs are remarkable for their wide range of optical colors, ranging from nearly neutral to very red. In fact, two of the most famous Centaurs span the full range of colors present in the outer solar system: (2060) Chiron is nearly neutral (Hartmann et al. 1990; Meech & Belton 1990; Luu & Jewitt 1990) and shows weak bands of water ice (Foster et al. 1999; Luu et al. 2000) while (5145) Pholus is very red (Mueller et al. 1992), perhaps because of the presence of irradiated complex organic compounds (Luu et al. 1994; Cruikshank et al. 1998). The fact that Chiron is an active Centaur while Pholus is not has led several investigators to speculate that outgassing activity might be responsible for the neutral colors of some Centaurs. For example, outgassing could blanket the surface of a body with neutral-colored fine particles excavated from beneath the few meter-thick surface layer in which radiation damage is strong (Delsanti et al. 2004). There is some evidence that the Centaurs as a whole occupy a bimodal distribution of optical colors (Peixinho et al. 2003). If so, the bimodality could tell us which Centaurs have been recently active. However, other interpretations are possible. Perhaps

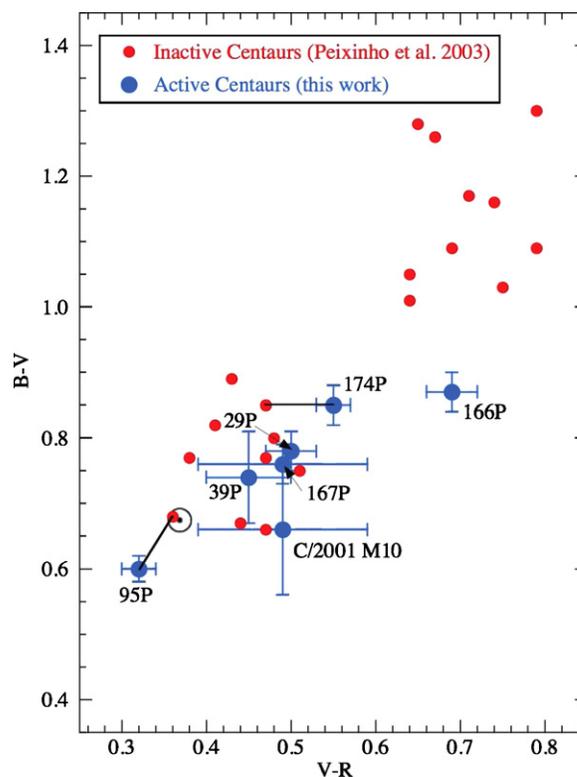


Figure 9. $B - V$ color index is plotted against the $V - R$ index, using data from Table 6 for the active Centaurs (blue circles) and from Peixinho et al. 2003 for the inactive Centaurs (red circles). Objects present in both samples are connected by straight lines. The color of the Sun is marked.

(A color version of this figure is available in the online journal.)

activity does not affect the surface colors. The color differences could reflect different intrinsic compositions, possibly related to different Centaur source regions.

Figure 9 shows the distribution of the Centaurs in the $B - V$ versus $V - R$ color plane. The data on inactive Centaurs (red circles) are taken from the compilation of Peixinho et al. (2003), while data on active Centaurs (blue circles with error bars) are from Table 6. Measurements of two Centaurs (Chiron and 60558) appearing in both the Peixinho et al. compilation and the present work are joined by lines. For clarity, only the uncertainties on the photometry of the active Centaurs are shown. The Figure shows the “red” and “blue” color clumps occupied by the Centaurs in the Peixinho sample. Six of the active Centaurs for which we possess quality $B - V$ and $V - R$ colors fall within the “blue” clump, while the seventh, 166P, is red in $V - R$ but not in $B - V$, and falls in neither the blue clump nor the red.

An observation relevant to the more neutral colors of the active Centaurs is provided by Figure 10, which shows the $B - I$ color index measured as a function of the projected angular distance from the photocenter of the five active Centaurs. Only these five objects have adequate spatial data and signal-to-noise ratio to justify an examination of the radial color gradient. The wide range in the $B - I$ colors of the active Centaurs is apparent from this Figure, from the very red 166P/2001 T4 to nearly neutral C/2001 M10 (the Solar color index is $(B - I)_\odot = 1.38$). The main trend apparent in Figure 10 is the tendency for the $B - I$ color index to decrease with angular radius, meaning that the comae of the active Centaurs are typically bluer than their centers. We used the profiles of field stars to confirm that the

Table 6
Adopted Colors of the Active Centaurs

Object	UT 2005	ϕ^a	R^b	$B - V^c$	$V - R^c$	$R - I^c$
29P/SW1 (coma)	2002 Sep 08	2.2	16.32 ± 0.02	0.78 ± 0.03	0.50 ± 0.03	0.52 ± 0.03
39P/Oterma (coma)	2002 Sep 07	2.2	21.83 ± 0.05	0.74 ± 0.07	0.45 ± 0.05	0.59 ± 0.10
166P/2001 T4 (coma)	2001 Nov 18	2.2	20.01 ± 0.02	0.87 ± 0.03	0.73 ± 0.03	0.71 ± 0.03
(2060) 95P/Chiron ^d	2001 Jun 13	1.4	15.77 ± 0.04	0.60 ± 0.03	0.32 ± 0.03	0.64 ± 0.03
174P/Echeclus (60558) ^e	2006 Mar 23	1.0	20.13 ± 0.1	0.71 ± 0.15	0.56 ± 0.15	...
174P/Echeclus (60558) ^f	2006 Feb 24	30	15.18 ± 0.04	0.81 ± 0.07	0.50 ± 0.06	0.58 ± 0.05
C/2001 M10 (coma)	2004 Oct 10	2.2	22.08 ± 0.06	0.66 ± 0.10	0.51 ± 0.03	0.51 ± 0.03
167P/2004 PY42 (nucleus)	2004 Oct 10	2.2	20.69 ± 0.02	0.80 ± 0.03	0.49 ± 0.10	0.42 ± 0.08
P/2004 A1	2006 Oct 01	3.3	19.76 ± 0.02	...	0.32 ± 0.03	...
<i>Solar Colors</i>				<i>0.67</i>	<i>0.36</i>	<i>0.35</i>

Notes.

^a Radius of the circular aperture within which the measurement was taken, in arcseconds.

^b R -band magnitude.

^c Color index, as listed, with 1σ uncertainty.

^d Romon-Martin et al. 2003.

^e Rousselot 2008; this small aperture measurement samples the primary nucleus.

^f Bauer et al. (2008); this large aperture measurement samples the primary and secondary nuclei plus the associated dust envelope.

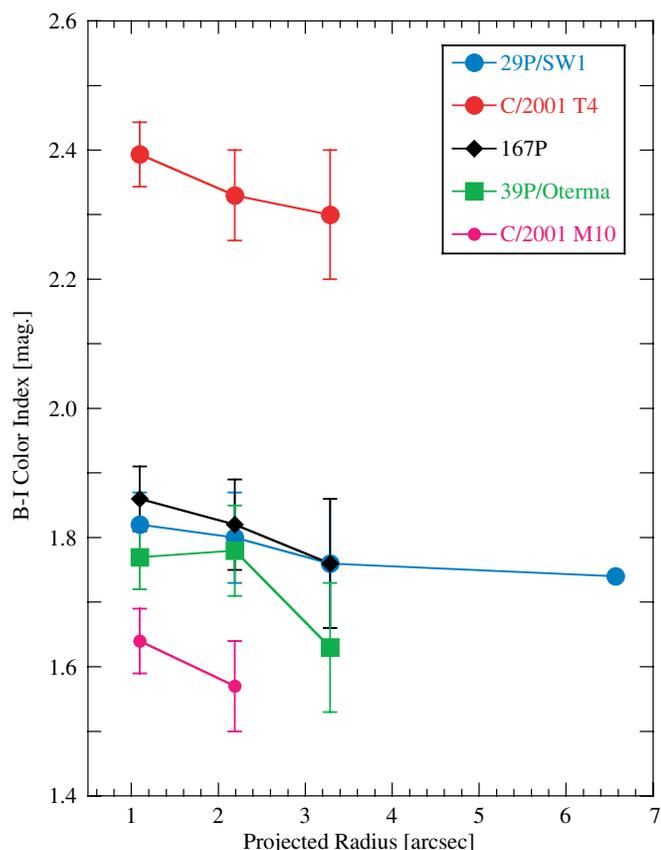


Figure 10. Average $B - I$ color index vs. projected radial distance from the photocenter for five active Centaurs. General trend toward bluer colors at larger radii is apparent.

(A color version of this figure is available in the online journal.)

color gradients in Figure 10 could not be artifacts produced by differences in the image quality between the B and I filters. A radial color gradient in the observed sense has been observed for some active JFCs. For example, P/Tempel 2 had quite different nucleus and coma colors of $V - R = 0.55 \pm 0.03$ and 0.36 ± 0.03 , respectively (A'Hearn et al. 1989; Jewitt & Luu 1989). Relatively blue coma is plausibly interpreted as an

effect of nongeometric scattering in optically small particles, so perhaps Figure 10 simply shows that the particles ejected from the active Centaurs are, on average, small enough for geometric effects to imbue them with a color different from the underlying nucleus material. More complex spatial gradients have been previously reported in 29P/SW1, where they may be associated with particle size sorting by solar radiation pressure (Cochran et al. 1982). Unfortunately, no detailed study of color gradients in cometary comae has been published, and the relation between the coma color and the color of the underlying nucleus cannot be specified with confidence.

In summary, the active Centaurs tend to lack the ultrared matter observed in the Kuiper Belt and some inactive Centaurs. We find no evidence that the colors of the active Centaurs are bimodally distributed (Figure 9) but the number of well-measured objects is small. The colors of the active Centaurs might reflect a true lack of ultrared matter on their nuclei, or simply an artifact of color dilution with neutral-blue coma particles in the near-nucleus environment (Figure 10). In view of the complications imposed by coma contamination of the near-nucleus photometry, we are reluctant to draw any strong conclusions based on the correlation between Centaur color and activity.

3.4. Orbital Distribution of the Active Centaurs

Figure 11 shows the distribution of the Centaurs in the semimajor axis, a , versus eccentricity, e , plane. Active Centaurs, including those observed in the present sample plus others reported in the literature as being active, are shown as blue circles, while all other Centaurs are shown in red. Also plotted (green circles) are the JFCs defined as having a Tisserand parameter with respect to Jupiter $2 \leq T_J \leq 3$. Lines are plotted in Figure 11 to mark the loci of orbits having perihelion distances equal to the semimajor axes of the giant planets. Objects above the line marked " q_J ," for instance, have perihelion inside Jupiter's orbit.

Figure 11 shows that the distribution of the active Centaurs in the a versus e plane is different from that of the Centaurs as a whole. The active Centaurs tend to possess smaller than average perihelia, with most located in a band between the orbits of Jupiter (curve labeled q_J in Figure 11) and Saturn (q_S). The

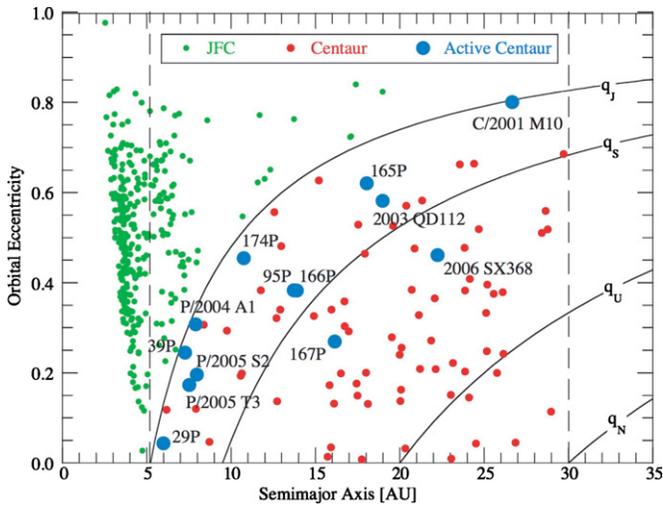


Figure 11. Orbital semimajor axis vs. eccentricity for active (blue) and inactive (red) Centaurs, and for JFCs (green). Labeled curves show the loci of orbits having perihelia equal to the semimajor axes of the four giant planets. Vertical dashed lines show the semimajor axes of Jupiter and Neptune, bounding the region of the Centaurs.
(A color version of this figure is available in the online journal.)

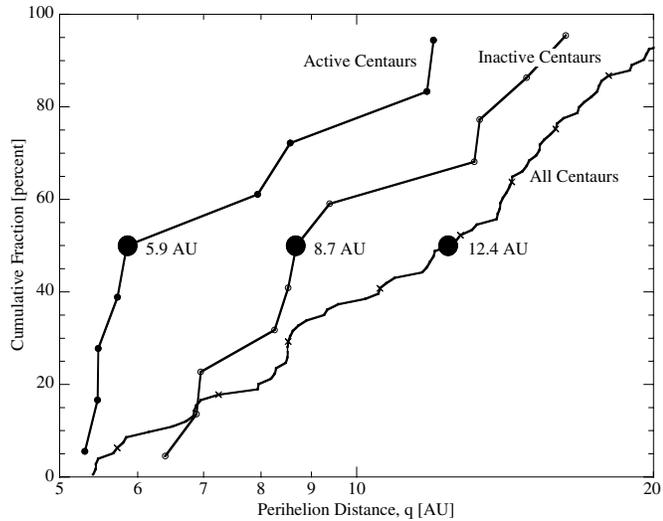


Figure 12. The cumulative distributions of the perihelion distances of, from left to right, the active Centaurs in our sample, the inactive Centaurs in our sample and all known Centaurs. The median perihelion distances for each group are marked.

median (mean) perihelion distance of the active Centaurs in our sample is 5.9 AU (7.6 ± 0.9 AU, $N = 9$ data points) whereas the corresponding quantities for the inactive Centaurs in our sample are 8.7 AU (10.2 ± 1.1 AU, $N = 11$). The Centaurs as a whole have median (mean) perihelion distances 12.4 AU (12.5 ± 0.5 AU, $N = 89$). The difference is emphasized in Figure 12, which shows the cumulative distributions of the perihelion distance for the active and inactive Centaurs and for the total Centaur population. According to the Kolmogorov–Smirnov test, the statistical likelihood that the active and inactive Centaur perihelion distance distributions could be drawn by chance from the same parent population is 0.03, corresponding to about a 2.2σ significance. The corresponding numbers for the active Centaurs and the total Centaur population are 0.01 and $\sim 2.7\sigma$.

Could these differences be artifacts of an observational bias? For example, is it possible that Centaurs of all perihelion distances are active but that the mass loss can be detected only in those with the smallest q ? To investigate this possibility, we must understand the relation between the detection of a coma and the rate of mass loss from the nucleus responsible for that coma.

3.5. Mass-Loss Rates

We seek to relate the optical photometry to the mass production rates of the Centaurs. To do this, we employ a three-step procedure. First, the photometry is used to estimate the scattering cross-section of the dust in the coma of each object. Second, the cross-section is related to the dust mass through the adoption of a size distribution of dust particles that is based on independent measurements from comets near 1 AU. Third, the mass-loss rate is inferred by dividing the dust mass by the time of residence of the dust grains as they travel across the projected photometry aperture. There are many unknowns in each step of this procedure and the derived mass-loss rates are, at best, crude estimates of the true values. Nevertheless, in the absence of more information, this is the best we can do.

Magnitude m_d (Equation 1) is related to the geometric albedo, p_R , and total cross-section of the coma dust particles within the projected annulus, C_d , by the inverse square law, conventionally expressed (Russell 1916) as

$$p_R \Phi(\alpha) C_d = 2.25 \times 10^{22} \pi R_{\text{au}}^2 \Delta_{\text{au}}^2 10^{-0.4(m_d - m_\odot)}. \quad (3)$$

Here, R_{au} and Δ_{au} are the heliocentric and geocentric distances expressed in AU, α is the phase angle and m_\odot is the apparent magnitude of the Sun viewed from 1 AU. The quantity $\Phi(\alpha)$ represents the phase darkening. We further take $p_R = 0.1$, consistent with measurements of active comets in backscattering geometry (Kolokolova et al. 2004). Equations 1 and 3 together give the dust cross-section for each Centaur, which we list in Table 4. The dust cross-section and the dust mass are related through the size distribution and other parameters of the dust (see Appendix). We employ Equation (A4) to relate C_d , the cross-section derived from the photometry to M_d , the dust mass. Values of M_d are also listed in Table 4.

The time of residence of the dust grains, at heliocentric distance R_{au} , in the annulus defined by ϕ_1 and ϕ_2 is just

$$\tau(R_{\text{au}}) = 1.5 \times 10^{11} \Delta_{\text{au}} (\phi_2 - \phi_1) / v(R_{\text{au}}), \quad (4)$$

where ϕ_1 and ϕ_2 are expressed in radians, and $v(R_{\text{au}})$ is the radial outflow speed of the dust grains from the nucleus. The heliocentric variation of the dust velocity is uncertain but is known to decrease with increasing heliocentric distance, albeit slowly. For example, measurements of expansion speeds in C/Hale–Bopp over the distance range $4 \leq R_{\text{au}} \leq 14$ (Biver et al. 2002) are compatible with

$$v(R_{\text{au}}) = v_0 \left[\frac{R_0}{R_{\text{au}}} \right]^{1/4} \quad (5)$$

where $v_0 = 550 \text{ m s}^{-1}$ is the velocity at reference distance $R_0 = 5$ AU. Equation (5) likely gives an upper limit to the true speed, since the Hale–Bopp measurements refer to gas. Entrained dust will, in general, not be perfectly coupled to the gas and therefore should travel more slowly.

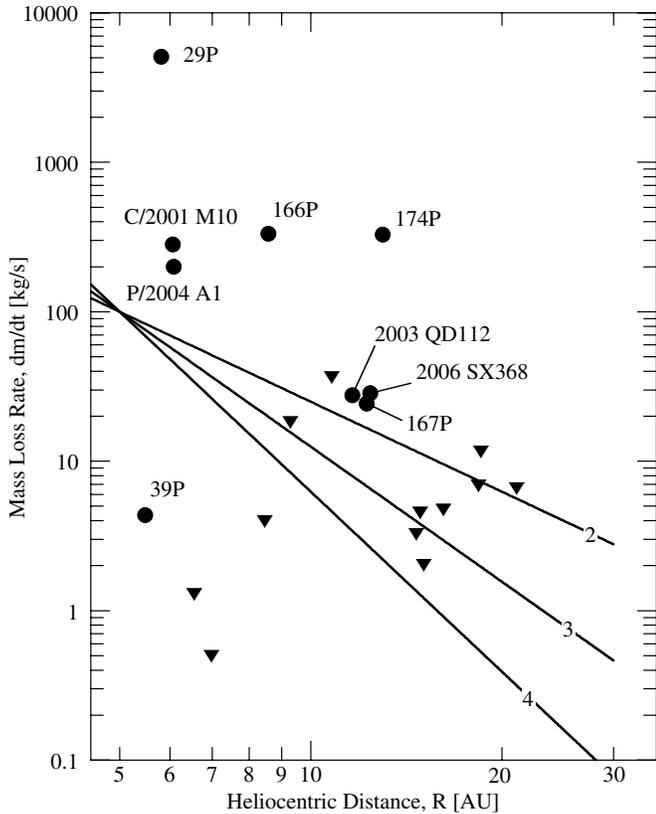


Figure 13. Derived dust mass-loss rates vs. heliocentric distance for the Centaurs in this sample. Circles represent measurements of the active Centaurs while inverted triangles represent upper limits to dust on Centaurs appearing inactive. The solid lines show power-laws of slope $n = 2, 3$ and 4 , as marked, for reference (see Equation (7)).

From Equations (4) and (5), the time taken for dust grains traveling radially to cross the annulus is written

$$\tau(R_{\text{au}}) = 1.5 \times 10^{11} \frac{\Delta_{\text{au}}(\phi_2 - \phi_1)}{v_0} \left[\frac{R_{\text{au}}}{R_0} \right]^{1/4}. \quad (6)$$

This gives a simplistic estimate both because the grains do not radially cross the projected photometry annulus and because the grains presumably occupy a broad distribution of sizes, with the larger grains less efficiently coupled to the expanding gases of the coma and so traveling more slowly than small grains. Still, as a first approximation, Equation (6) provides a useful estimate. Column 7 of Table 4 shows that, for our data, values of τ typically fall in the range $(1 \text{ to } 4) \times 10^4$ s, or a fraction of a day. Model mass-loss rates computed from Equations (A4) and (6) are listed in Table 4.

The derived mass-loss rates are plotted against heliocentric distance in Figure 13, with filled circles to represent the active Centaurs and inverted triangles marking 3σ limits to the mass loss for those Centaurs in which no coma was detected. In passing, we note that the mass-loss rates for 29P estimated from our dust measurements ($\sim 5000 \text{ kg s}^{-1}$) are of the same order as rates estimated independently, although at different times, from measurements of gas in this object (~ 1000 to 2000 kg s^{-1} ; Senay & Jewitt 1994; Crovisier et al. 1995). Mass-loss rates from P/2004 A1 have also been independently estimated from imaging data by Mazzotta et al. (2006). Their estimates ($\sim 100 \text{ kg s}^{-1}$) are within a factor of two of those in Table 4 and Figure 13 ($\sim 200 \text{ kg s}^{-1}$). We are aware of no other mass-loss rate estimates for the active Centaurs, but conclude from

these limited comparisons that our simple photometric method produces values in reasonable accord with others.

The solid lines in Figure 13 show mass-loss rates varying as

$$\frac{dm}{dt} = 100 \left[\frac{5}{R_{\text{au}}} \right]^n \quad (7)$$

where index $n = 2, 3$ and 4 and the normalization to 100 kg s^{-1} at $R_{\text{au}} = 5$ is arbitrary. The Figure shows that those Centaurs observed to be active at $R_{\text{au}} < 13$ would also be measurably active at larger distances, if their sublimation followed an $n = 2$ dependence on R_{au} . For example 166P, when displaced from 8 AU to 20 AU, would lose mass at about 60 kg s^{-1} , according to the Figure, which is an order of magnitude larger than the detection limits established for other Centaurs (e.g., 52975; see Table 4) at this distance. (The one exception is the very weakly active 39P/Oterma: if displaced from 5.5 AU to larger distances an equilibrium coma around this object would quickly become unobservable.) Therefore, we conclude that the mass-loss rates in active Centaurs drop with heliocentric distance faster than R_{au}^{-2} . This is significant because $n = 2$ is the index expected from the equilibrium sublimation of an exposed supervolatile ice (because essentially all the incident Solar energy is used for sublimation, leaving none for thermal re-radiation). Solid carbon monoxide (CO), for instance, would sublimate in proportion to R_{au}^{-2} if exposed to sunlight. This suggests that the absence of coma around Centaurs with $R_{\text{au}} > 13$ is not an artifact of observational selection acting upon a set of Centaurs losing mass in proportion to R_{au}^{-2} . A radial variation with $n \geq 3$ can fit the data in Figure 13 but this implies that the mass loss is not driven simply by the sublimation of an exposed supervolatile. It is interesting to note that observations of comet C/Hale–Bopp as it moved out through the giant planet region of the solar system lend independent support to this conclusion. The mass-loss rate from the comet closely followed R_{au}^{-2} (Jewitt et al. 2008) and, indeed, coma has been consistently observed in optical images out as far as 26 AU (Szabó et al. 2008).

On this basis we conclude that the activity in the Centaurs is not driven by the equilibrium sublimation of CO or another supervolatile. Either the mass loss is driven by a less volatile material (which would have $n > 2$ in the observed range of heliocentric distances) or the restriction of the activity to smaller heliocentric distances is the result of a trigger which prevents mass loss when far from the Sun.

4. DISCUSSION

4.1. Origin of the Activity

The prime observational clues to the origin of Centaur activity are:

1. *Perihelion distribution.* The median perihelion distance of the active Centaurs is 5.9 AU, whereas the median perihelion distance for the Centaurs as a whole is 12.4 AU (Figure 12). As discussed above (see Figure 13), it is unlikely that this difference is the result of observational selection acting on a set of active Centaurs in the strong sublimation limit. Instead, the data suggest that the activity is triggered or driven by a temperature-related process.
2. *Detection of CO.* Centaur 29P is a strong source of carbon monoxide, with a sustained outgassing rate 1000 kg s^{-1} to 2000 kg s^{-1} (Senay & Jewitt 1994; Crovisier et al. 1995; Gunnarsson et al. 2008). CO has also been reported (but not confirmed) with a similar outgassing rate in Centaur (2060)

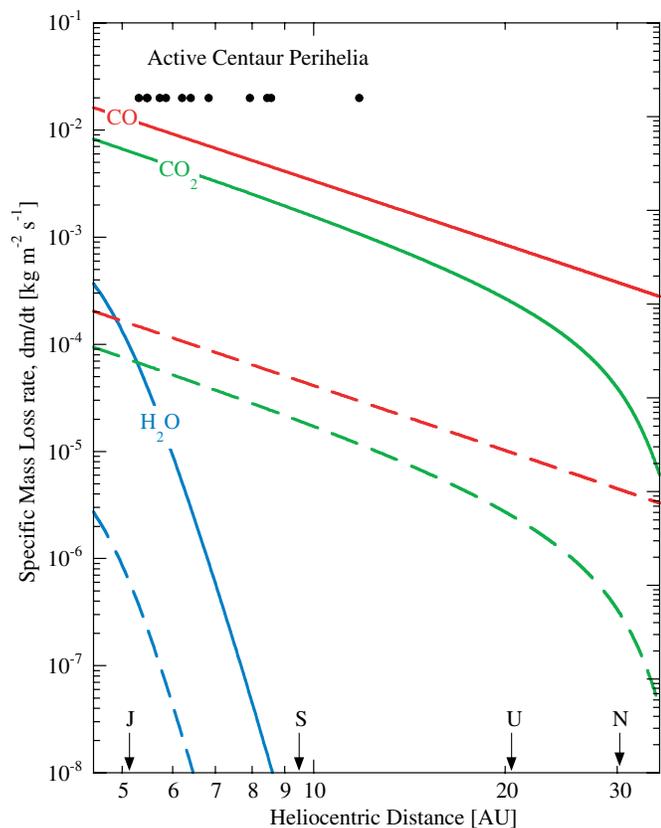


Figure 14. Left axis: specific mass-loss rate ($\text{kg m}^{-2} \text{s}^{-1}$) as a function of the heliocentric distance (AU) computed from Equation (8). Blue curves apply to crystalline water ice, green curves to CO_2 ice and red curves to CO ice. Solid and dashed lines for each ice refer to the high and low temperature limiting cases, as described in Section 4.2. The radii of the orbits of the giant planets are marked for reference at the bottom of the plot. The perihelion distances of the active Centaurs are marked at the top with black circles.

(A color version of this figure is available in the online journal.)

Chiron (Womack & Stern 1999). While other Centaurs show no spectral evidence for CO sublimation or, indeed, for any other volatile (Bockelée-Morvan et al. 2001; Jewitt et al. 2008), these detections focus attention on the likely role of this supervolatile in driving mass loss from the active Centaurs.

3. *Line Profiles of CO.* The rotational line profiles of CO in 29P are asymmetric, with a narrow, blueshifted component (Senay & Jewitt 1994; Crovisier et al. 1995; Gunnarsson et al. 2008). This shows that the release of the CO is diurnally modulated and therefore that the CO source is in good thermal contact with (i.e., close to) the physical surface of 29P. The source cannot be deep.

We now expand upon these points.

Figure 14 shows solutions to the energy balance equation

$$\frac{L_{\odot}(1-A)}{4\pi R^2} = \chi \left[\epsilon\sigma T^4 + L(T) \frac{dm}{dt} \right] \quad (8)$$

in which $L_{\odot} = 4 \times 10^{26} \text{ W}$ is the luminosity of the Sun, R is the heliocentric distance in meters, A is the Bond albedo, ϵ is the emissivity, and $\sigma = 5.67 \times 10^{-8} \text{ W m}^{-2} \text{ K}^{-4}$ is the Stefan–Boltzmann Constant. The second term on the right-hand side accounts for sublimation at the rate dm/dt ($\text{kg m}^{-2} \text{ s}^{-1}$), and $L(T)$ (J kg^{-1}) is the (temperature-dependent) latent heat of sublimation. Conduction is neglected from Equation (8)

based on the extremely small thermal inertias of two measured Centaurs (Fernández et al. 2002) and of the nucleus of P/Tempel 1 (Groussin et al. 2007). Quantity χ is a dimensionless parameter intended to represent the way in which the energy from the Sun is distributed over the surface of the body and is affected by both the thermal parameters of the nucleus and its rotational state. Physically, χ may be interpreted as the ratio of the area from which absorbed sunlight is radiatively lost into space to the area over which energy is absorbed. We considered two limiting cases to bracket the likely range of χ at any given heliocentric distance. A flat surface exposed normally to the Sun absorbs and radiates from the same area and so has $\chi = 1$. This limit would apply, for example, to the subsolar point of a slowly rotating body or to any nucleus whose projected pole direction passed through the Sun. At the other extreme, an isothermal sphere of radius r absorbs over πr^2 but radiates from $4\pi r^2$, and so has $\chi = 4$. We took $L(T)$ from Brown & Ziegler (1979) and assumed $A = 0.1$, $\epsilon = 0.9$ (the particular values assumed do not have a critical effect on the computed temperatures or mass-loss rates provided $A \ll 1$ and $\epsilon \gg 0$).

The solutions to Equation (8) are shown in Figure 14 for crystalline water ice (blue lines), CO_2 ice (green lines) and CO ice (red lines). For each ice we plot two curves to show the low and high χ values. Furthermore, at each heliocentric distance we calculated the critical grain size, a_c , above which the gas drag force (proportional to dm/dt) is insufficient to launch a grain against the gravitational attraction to a nucleus of density $\rho = 10^3 \text{ kg m}^{-3}$ and radius $r_n = 25 \text{ km}$. Figure 14 shows solutions of Equation 8 for which $a_c > 1 \mu\text{m}$. Grains much smaller than $1 \mu\text{m}$ are inefficient optical scatterers with characteristically blue colors that contradict the measured colors of the Centaurs. Figure 14 shows that CO and CO_2 sublimation is able to eject $1 \mu\text{m}$ grains over the full range of heliocentric distances swept by the Centaurs while H_2O sublimation can only do so for $R_{\text{au}} \leq 6.5$ and 8.5 , in the low and high temperature limits, respectively.

Figure 14 shows that none of the ices provide a convincing match to the Centaur activity. The Centaurs are too cold, even at their subsolar points, for water sublimation to be effective except close to Jupiter’s orbit. For example, 2003 QD112, 166P, 167P, and 2006 SX368 are active at 11.6, 8.6, 12.2, and 12.4 AU, all far beyond the water sublimation zone. Other Centaurs such as C/2001 M10, P/2004 A1, 29P are active at distances (6.06–9.2 AU, 6.1 and 5.8 AU, respectively) where water ice can weakly sublimate, but a comparison of the specific mass-loss rates (Figure 14) with the actual mass-loss rates (Table 4) shows that implausibly large areas of exposed water ice are required. More importantly, spectroscopic observations clearly show that the mass loss is controlled by large fluxes of CO, at least for 29P (Senay & Jewitt 1994; Crovisier et al. 1995), not water ice.

Conversely, CO and CO_2 ices are so volatile that they should sublimate strongly throughout the planetary region. If free CO ice were present on the Kuiper Belt objects and Centaurs we should observe comae around these bodies at all heliocentric distances out to, and beyond, the orbit of Neptune. Since this is clearly not the case, we conclude that free CO is not responsible for the observed activity. Nondetections of CO rotational lines have been used to place empirical limits to the possible fractional areas of exposed CO on Centaurs and KBOs near 0.1%–1%, implying a strong depletion (Jewitt et al. 2008). Physically, this conclusion is easy to understand. Carbon monoxide ice sublimates so rapidly that any exposed CO could survive only briefly. For example, even in the cold limit, CO sublimates at 30 AU at a rate near $dm/dt = 10^{-5} \text{ kg m}^{-2} \text{ s}^{-1}$

(Figure 14), corresponding to recession of the ice surface (for density $\rho = 1000 \text{ kg m}^{-3}$) at 10^{-8} m s^{-1} . In one orbital period of ~ 150 yrs, 50 meters of CO would sublimate away. In 1000 orbits (only ~ 150 kyr) the sublimation distance would rival the radius of the Centaur. Clearly, free CO ice cannot persist near the surface of any Centaur or Kuiper Belt object and might only be found, if anywhere, at the bottom of deep, self-shadowing pits created by sublimation. The only exception to this statement applies to the largest Kuiper Belt objects, such as Pluto, on which sublimated CO can be recycled through a thin, gravitationally bound atmosphere. The much smaller objects that dominate the Centaur population must all be free of exposed supervolatile ices on their surfaces.

Could the CO be buried beneath a refractory layer and heated only slightly by the Sun? In this way, mass loss at large R_{au} might be strongly suppressed relative to the $n = 2$ heliocentric dependence expected of supervolatile ice. The main argument against this possibility is provided by the profiles of the CO rotational lines recorded in submillimeter spectra. The lines show a narrow, blueshifted component that indicates that CO is released from the source primarily on the sun-facing side. Given that all small bodies are spinning, with characteristic periods typically of order $t_s \sim 10$ hr, the CO line asymmetry requires that the CO be very close to the physical surface. To see this, consider a material in which the thermal diffusivity is $\kappa = 10^{-7} \text{ m}^2 \text{ s}^{-1}$, as representative of a finely powdered dielectric. The diurnal thermal wave has a characteristic scale $\ell \sim (\kappa t_s)^{1/2}$. Substituting, we estimate $\ell \sim 0.06 \text{ m}$ (6 cm). If the CO is buried any deeper than ℓ , the heating of the CO will be unable to respond to the day-night insolation cycle, and the CO line profile will not show the sunward peak that is observed. This argument seems robust in 29P and proves that the CO is very close to the surface of this nucleus. However, we need more line profile measurements in other Centaurs in order to be sure that sunward outgassing is a general property of these objects.

4.2. Amorphous Ice in the Centaurs

Amorphous water ice is thermodynamically unstable and converts exothermically to the crystalline form at a rate that depends on the ice temperature. The energy released per unit mass of pure water ice is about $9 \times 10^4 \text{ J kg}^{-1}$, and the timescale for the conversion is given in years by

$$\tau_{\text{cr}} = 3.0 \times 10^{-21} \exp \left[\frac{E_A}{kT} \right] \quad (9)$$

where E_A is the activation energy, k is Boltzmann's constant, and $E_A/k = 5370 \text{ K}$ (Schmitt et al. 1989). Setting $\tau_{\text{cr}} = 1 \text{ hr}$ ($\sim 10^{-4} \text{ yr}$), to correspond to the timescale of a typical laboratory experiment, the crystallization temperature is given by Equation (9) as $T_{\text{cr}} = 140 \text{ K}$. In solar system bodies where the timescales are much longer, significant crystallization can occur at correspondingly lower ice temperatures. At temperatures $T < 77 \text{ K}$, the crystallization time exceeds the age of the solar system. Amorphous ice is highly porous and can contain within its structure a large number of atoms and molecules trapped at the time of condensation of the ice from vapor. Experiments show that the trapped component is released upon crystallization, as the amorphous structure rearranges into a crystalline lattice and the trapped molecules are squeezed out (Notesco & Bar-Nun 1996; Notesco et al. 2003). On laboratory timescales this happens at $\sim 140 \text{ K}$ but in solar system bodies the gas release

would occur more slowly at lower temperatures. The timescale given in Equation (9) is subject to uncertainties of extrapolation (the timescale is measured at temperatures much higher than relevant in the outer solar system) and might be very different for ice formed in the presence of contaminants or particles, like silicate dust, that could act as crystallization nuclei.

The possible role of crystallization in the activity of comets has been studied in great detail (e.g., Enzian et al. 1997; Klinger et al. 1996; Notesco & Bar-Nun 1996; Notesco et al. 2003; Prialnik 1997). By providing an extra energy source, crystallization enables these models to be broadly successful in explaining mass loss from comets at temperatures too low (and heliocentric distances too large) for crystalline water ice to sublimate. The crystallization models can be criticized on several levels. Disconcertingly, there is no direct evidence for the existence of amorphous ice in cometary nuclei (however, in fairness we note that observations provided no direct evidence for water ice in the comets for many decades after the formulation of Whipple's (1950) dirty ice nucleus model). The results of crystallization models depend upon a large number of unknown or poorly constrained parameters (e.g., thermal diffusivity, porosity, tortuosity, bulk composition, optical parameters including albedo and emissivity, and rotational properties) for each of which values must be assumed in order to fit the data. Kouchi & Sironi (2001) have further argued that the transition in real ice might not be exothermic at all, because the sublimation of trapped volatiles liberated at crystallization would carry away all of the energy released. We feel that this objection has merit within a mean free path of the escape surface but, in buried ice, the volatiles cannot escape directly to space and the energy of the phase transition would remain trapped.

While observational evidence for amorphous ice in the Centaurs is, at best, indirect, evidence for water ice clathrate is entirely lacking. The clathrates trap guest molecules within the periodic lattice structure of crystalline water ice. Hence, the volatility of clathrate ice is essentially the same as that of unclathrated, crystalline water ice (Klinger et al. 1996). The perihelion distribution of the active Centaurs, which extends far beyond the range of distances over which crystalline water ice can sublimate (Figure 11) is therefore inconsistent with the sublimation of clathrated water ice. Furthermore, the conditions of low temperature and low pressure likely to be found in the protoplanetary disk of the Sun, where the Centaurs formed, are more conducive to the formation of amorphous forms than they are to clathrates (Devlin 2001). If clathrates exist in the Centaurs, they have nothing to do with the observed outgassing activity.

Here, we suggest a simple estimate of where in the solar system, if exposed amorphous ice is present, the phase transition may be important by comparing the crystallization time, τ_{cr} , to the orbital period, t_k . We reason that this gives a highly conservative criterion, because we can be sure that amorphous ice will have crystallized on objects for which $\tau_{\text{cr}} \ll t_k$ but crystallization may also have occurred on a longer timescale, up to a maximum set by the evolution of the orbit of the Centaur. In the spirit of the criterion, we consider circular orbits, for which $t_k = R_{\text{au}}^{3/2}$ (yr).

We bracket the range of likely surface temperatures as in Section (5) using Equation (8) with $dm/dt = 0$. The highest surface temperature should be that at the subsolar point, for which $\chi = 1$, giving $T(R_{\text{au}}) = T_1/R_{\text{au}}^{1/2}$ with $T_1 = 392 \text{ K}$. The lowest surface temperature is attained using the isothermal blackbody approximation, for which $\chi = 4$ and $T_1 = 280 \text{ K}$. Then, from Equation (9) the critical distance inside which

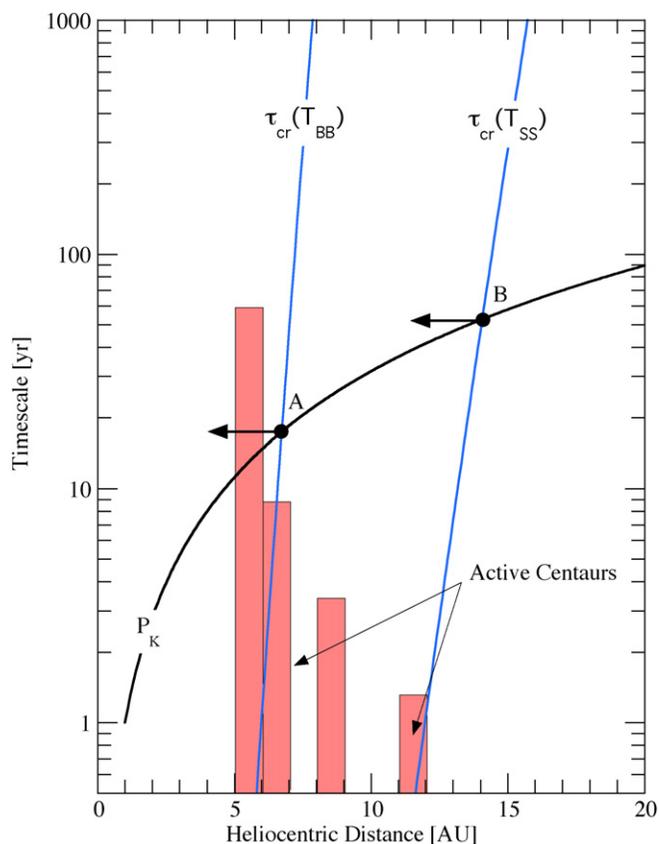


Figure 15. The black line marked P_K shows the variation of the Keplerian period with semimajor axis. Blue lines show the crystallization timescales for the lowest and highest likely equilibrium temperatures, as described in Section 4.2. The crystallization and orbital periods are equal at the points labeled *A* and *B*, having heliocentric distances of ~ 6.5 AU and 14 AU, respectively, computed from Equation (10). Left-pointing arrows indicate that crystallization is expected at all smaller heliocentric distances. The histogram shows the distribution of the perihelion distances of the active Centaurs.

(A color version of this figure is available in the online journal.)

exposed ice must be crystalline is found from the solution to

$$\frac{3}{2} \ln R_{\text{au}} = -47.26 + \frac{E_A \sqrt{R_{\text{au}}}}{kT_1} \quad (10)$$

Solving for $280 < T_1 < 392$ K gives $6.8 < R_{\text{au}} < 14.0$ for the distances inside which the phase transition is expected (see also Figure 15). All of the active Centaurs have perihelia small enough for the crystallization of amorphous water ice to be a contributing factor. This observation is suggestive, albeit indirect, evidence for the existence of amorphous ice in the Centaurs.

There are obvious difficulties in translating the results of laboratory experiments conducted typically on timescales of hours to the relevant cometary timescales, measured in millions of years and longer. Still, the overall picture is very attractive. The Centaur precursors in the Kuiper Belt contain amorphous ice, laden with trapped molecules including CO. Once deflected into the planetary region, a thermal wave of growing amplitude begins to propagate beneath the surface. Regions within the Centaurs that attain peak temperatures above T_{cr} begin to convert to the crystalline phase, releasing trapped gases and the energy of the transition. This occurs preferentially on the small-perihelion Centaurs but post-perihelion activity can be expected because of the finite timescales for the transport of heat in porous media.

Most of the active Centaurs were discovered soon before perihelion and were active at discovery. This is shown in Table 7, where we list the date of perihelion, the date of the first reported activity, the lag between these two dates, ΔT , the orbital period, P , and the ratio $\Delta T/P$. The interval between these dates, ΔT , is typically measured in months, only a small fraction of the orbital period, also listed in Table 7. Discovery near perihelion is an artifact of observational bias in magnitude limited surveys, and holds no significance for us here. The observation that mass loss is underway close to and even before perihelion is more interesting. It tells us that there is in general no great thermal lag between the period of maximum heating and the onset of mass loss. The thermally driven process behind the production of coma is located close to the surface, just as we concluded earlier from the blueshifted CO emission lines observed in 29P.

Observational evidence relating to the physical state of ice on the Centaurs is limited. (10199) Chariklo (formerly 1997 CU26) shows water absorptions at $1.5 \mu\text{m}$ and $2.0 \mu\text{m}$ but there is no evidence for the $1.65 \mu\text{m}$ feature that is characteristic of crystalline ice (Dotto et al. 2003). Given the quality of the spectrum and the muted depths of the $1.5 \mu\text{m}$ and $2.0 \mu\text{m}$ features, though, the absence of the $1.65 \mu\text{m}$ band cannot be used as evidence for amorphous ice. (2060) 95P/Chiron likewise shows water ice absorptions (Foster et al. 1999; Luu et al. 2000) but with absorption band strengths that are modulated by the activity of the coma. The bands are again too weak and the data are too poor to permit us to rule definitively on the crystalline versus amorphous state of the ice. The activity in 2060 Chiron has been modeled as a product of crystallization heating by Prrialnik et al. (1995).

4.3. Amorphous Ice in the Comets and Kuiper Belt Objects

The available data are consistent with the presence of amorphous ice in the Centaurs. What does this mean for the dynamically related populations of JFCs and Kuiper Belt objects?

A fraction of the Centaurs will eventually be captured as JFCs, suggesting that these bodies should also contain amorphous ice. Observationally, direct measurements of the ice in comets are very few, and the crystalline state is not well measured. However, as noted above, crystallization has long been invoked by modelers as an important process in the JFCs, with an apparent success that is tempered by knowledge of the large number of poorly constrained parameters that must be assumed. Crystallization models also provide plausible fits to the activity of some long-period and dynamically new (Oort cloud) comets (Meech et al. 2009). A fair statement would be that there is little evidence concerning the crystallinity of cometary water ice and nothing to strongly contradict the presence of ice in the amorphous state (Davies et al. 1997).

A somewhat different circumstance prevails for the precursors to the Centaurs located in the Kuiper Belt. There, near infrared spectral observations provide a seemingly contradictory result; the $1.65 \mu\text{m}$ absorption band which is diagnostic of the presence of crystalline water ice is clearly observed in many KBOs (Jewitt & Luu 2004; Trujillo et al. 2007). Some fraction of the ice could be amorphous and still fit the measured profile of the $1.65 \mu\text{m}$ band (Merlin et al. 2007; Pinilla-Alonso et al. 2009). On the other hand, there are no well-observed objects in which the characteristic $1.5 \mu\text{m}$ and $2.0 \mu\text{m}$ bands of water ice are observed but the $1.65 \mu\text{m}$ crystalline ice band is absent. Therefore, where the data are good enough to make a diagnosis, the available spectra provide no evidence that amorphous is ubiquitous on the KBOs. If the Centaurs contain

Table 7
Onset of Activity

Object	Perihelion Date ^a	Activity Date ^b	ΔT (yr) ^c	P (yr) ^d	$\frac{\Delta T}{P}$	Comment
C/2001 M10	2001 Jun 21	2001 Jun 20	0.0	137.7	0.00	Active at discovery
P/2004A1(LONEOS)	2004 Aug 25	2004 Jan 13	−0.6	22.2	−0.03	Active at discovery
39P/Oterma	2002 Dec 21	2002 Sep 07	−0.3	19.5	−0.02	...
29P/SW1	2004 Jun 30	14.6	...	Always active
174P/Echeclus (60558)	2015 May 06	2006 Jan	−9.4	35.2	−0.27	...
P/2005T3 (Read)	2006 Jan 13	2005 Oct	−0.3	20.6	−0.01	Active at discovery
P/2005 S2 (Skiff)	2006 Jun 30	2005 Oct	−0.8	22.5	−0.04	Active at discovery
165P/Linear	2000 Jun 15	2000 Jan 29	−0.5	76.6	−0.01	Active at discovery
2003 QD112	1998 May 29	2004 Oct 10	+4.4	82.7	+0.05	...
(2060) 95P/Chiron	1996 Feb 14	1989 Apr 10	−6.8	50.8	−0.13	Active since 1989
166P/2001 T4	2002 May 20	2001 Oct 15	−0.6	51.9	−0.01	Active since discovery
167P/2004 PY42	2001 Apr 10	2004 Aug 11	+3.3	64.8	+0.05	...

Notes.

^a Date of closest perihelion.

^b Dates of first reported cometary activity.

^c Time lag, in years, between perihelion and first date of reported activity. Negative means activity was detected before perihelion.

^d Orbital period, in years.

amorphous ice, and if they recently escaped from long-term storage in the Kuiper Belt, it is logical to conclude that ice in the KBOs should be, at least in part, amorphous. Certainly, amorphous ice could survive indefinitely in the KBOs given the very low radiation equilibrium temperatures in the Kuiper Belt. How might this seeming contradiction be explained?

One possibility is that this is a size-selection effect. The KBOs bright enough for the diagnostic 1.65 μm band to be spectroscopically detected tend to be large (~ 1000 km) objects. Heating from their gravitational binding energy (McKinnon 2002) or from trapped radioactive nuclei may have caused the ice in these large bodies to have crystallized, so explaining the spectroscopic data (Merk & Prialnik 2003). Smaller bodies of size comparable to the Centaurs listed in Table 3 (radii of a few $\times 10$ km or less) could have escaped crystallization but are effectively too faint to observe spectroscopically in the infrared. A second possibility is that the KBOs and Centaurs are stratified, with crystalline ice at the visible surface and amorphous ice beneath. For example, some process such as heating by micrometeorite bombardment could crystallize surface ice while leaving deeper ice in the amorphous state. Evidently, these two possibilities are not mutually exclusive.

4.4. The Future

A number of exciting observational prospects exist for making progress in this subject. First, the assertion that the active Centaurs are powered by the crystallization of amorphous ice is open to an observational test: this mechanism could be refuted if future observations reveal outgassing activity in Centaurs having perihelion distances much larger than the critical range for crystallization, as given by Equation (10). For example, the discovery of active Centaurs with perihelion $q > 20$ AU would be difficult to reconcile with the crystallization of amorphous ice, unless the dynamics showed that they had been very recently scattered outwards from an orbit with a smaller perihelion. Second, deep rotational line spectra with ALMA should show whether CO outgassing drives the mass loss in other active Centaurs, as is expected in the present hypothesis. Rotational line spectra having high sensitivity will also allow us to measure the mass-loss rates with greater accuracy, instead of relying on imperfect estimates based on observations of the dust, as

in this work. Third, measurements of the CO line profiles will show whether or not the mass loss is mainly sunward (and therefore close to the physical surface), as is observed in 29P/Schwassmann–Wachmann 1.

Future observations might also show why some Centaurs remain inactive even when their perihelia are small. In Figure 11, more than half of the Centaurs with perihelia between the orbits of Jupiter and Saturn appear inactive, although all are, presumably, hot enough to satisfy the crystallization criterion in Equation (10). This could be because the mass loss is intermittent or “bursty,” on timescales longer than the orbital period, as a result of instabilities in the crystallization front. Alternatively, these hot but inactive Centaurs might have been devolatilized in their near-surface regions as a result of past activity (i.e., they could have been in their current orbits for longer, as a group, than the active Centaurs). Still another possibility is that some Centaurs are deficient in volatiles, or perhaps lacking in the amorphous ice whose exothermic transformation appears to provide a plausible gas and energy source for the active Centaurs. Physical studies, careful long-term photometric monitoring and dynamical investigations of the Centaurs will help us to understand which of these possibilities might be correct.

5. SUMMARY

We present new observations of the active Centaurs, defined as solar system bodies whose orbits have both perihelia and semimajor axes between the orbits of Jupiter (at 5.2 AU) and Neptune (30 AU) and which are not in 1:1 resonance with the giant planets.

1. Including observations from the present work and from the literature, about 13% (12 of 92) Centaurs are measurably active, as shown by the presence of dust comae in optical images.
2. The active Centaurs have statistically smaller perihelion distances (median 5.9 AU) than the inactive Centaurs in our sample (median 8.7 AU) and smaller than the median perihelion distance of all known Centaurs (12.4 AU). Simulations show that the difference is unlikely to be an artifact of observational selection. Instead, the smaller distances of the active Centaurs appear to reflect the action of a thermally driven trigger or activity source.

3. The surface temperatures of the active Centaurs are too low (and the inferred mass-loss rates are too large) for the sublimation of crystalline water ice to constitute a plausible source of the activity.
4. Conversely, although gaseous carbon monoxide (CO) has been reported in two Centaurs (29P and Chiron), CO ice is too volatile to explain the measured perihelion distribution. If present, near-surface CO ice would strongly sublimate all the way out to Neptune's orbit, whereas no active Centaurs are known so far from the Sun.
5. The active Centaurs have orbital periods which are long compared to the timescale for the crystallization of water ice at their distance. No active Centaurs in which this inequality is reversed have yet been found. This fact is consistent with (but does not prove) crystallization as the trigger or driver of activity.
6. If Centaurs contain amorphous ice, as suggested by the available data, then their precursor bodies in the Kuiper Belt must likewise be at least partly amorphous and their progeny, the nuclei of JFCs, should also contain amorphous ice.
7. Most active Centaurs are slightly redder than the Sun at optical wavelengths, but they appear devoid of the ultrared matter found in the inactive Centaurs and Kuiper Belt objects. Physical interpretation of the colors, however, is stymied by coma contamination in the active Centaurs.

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APPENDIX

The connection between the cross-section of the particles and their mass depends largely upon the size distribution of the particles, but also on their density and scattering efficiency. For spherical particles of a single radius, a , the mass, M_d , and the cross-section, $C_d = \pi a^2$, are (in the geometric optics limit) proportional:

$$M_d = \frac{4}{3} \rho a C_d \quad (\text{A1})$$

where ρ is the grain bulk density.

This simple relation may be generalized for power-law size distributions in which the number of particles with radius between a and $a + da$ is $\Gamma a^{-q} da$, where Γ and q are constants of the distribution and the relation holds between minimum and maximum particle radii a_- and a_+ , respectively. The combined cross-section of such a distribution is

$$C_d = \pi \Gamma \int_{a_-}^{a_+} a^{2-q} da \quad (\text{A2})$$

while the total mass is

$$M_d = \frac{4}{3} \pi \Gamma \int_{a_-}^{a_+} a^{3-q} da. \quad (\text{A3})$$

Results obtained from the study of active comets show that the dust does not exactly follow a power law size distribution, but that a reasonable approximation is obtained by setting $q = 3.5$, $a_- = 0.1 \mu\text{m}$ and $a_+ = 1 \text{ cm}$, (e.g., Grün et al. 2001). Comparing Equations (A2) and (A3) then shows that the mass is given by

$$M_d = \frac{4}{3} \rho (a_- a_+)^{1/2} C_d \quad (\text{A4})$$

and this is the relation we employ here. The size of the average scatterer in this case is $\bar{a} = (a_- a_+)^{1/2} \sim 30 \mu\text{m}$, consistent with the observation that the comae of the Centaurs are neutral or reddish compared to the Sun, not blue (as would be expected if $\bar{a} \ll \lambda$, where $\lambda \sim 0.5 \mu\text{m}$ is the wavelength of observation).

Other relations can be trivially computed from Equations (A2) and (A3) and still more involved estimates can be obtained by including the wavelength- and size-dependent scattering efficiencies of the particles. However, given our basic ignorance of the properties of the dust in Centaurs, these extra steps hardly seem appropriate. In fact, uncertainties in the dust parameters set a fundamental limit to our ability to estimate dust mass production rates from optical data and the absolute values of the mass-loss rates derived from Equation (A4) could be wrong by an order of magnitude or more. In this paper, we use Equation (A4), with $(a_- a_+)^{1/2} = 30 \mu\text{m}$ and $\rho = 1000 \text{ kg m}^{-3}$, simply as a useful metric with which to compare relative mass-loss rates amongst the Centaurs.

REFERENCES

- A'Hearn, M. F., Campins, H., Schleicher, D. G., & Millis, R. L. 1989, *ApJ*, **347**, 1155
- Bauer, J. M., Choi, Y.-J., Weissman, P. R., Stansberry, J. A., Fernández, Y. R., Roe, H. G., Buratti, B. J., & Sung, H. I. 2008, *PASP*, **120**, 393
- Bauer, J. M., Fernández, Y. R., & Meech, K. J. 2003, *PASP*, **115**, 981
- Baum, W. A., Kreidl, T. J., & Schleicher, D. G. 1992, *AJ*, **104**, 1216
- Biver, N., et al. 2002, *Earth Moon Planets*, **90**, 5
- Bockelée-Morvan, D., Lellouch, E., Biver, N., Paubert, G., Bauer, J., Colom, P., & Lis, D. C. 2001, *A&A*, **377**, 343
- Bowell, E., Hapke, B., Domingue, D., Lumme, K., Peltoniemi, J., & Harris, A. W. 1989, in *Asteroids II* (Tucson, AZ: Univ. of Arizona Press), 524
- Brown, G., & Ziegler, W. 1979, *Adv. Cryog. Eng.*, **25**, 661
- Choi, Y.-J., Weissman, P. R., & Polishook, D. 2006, *IAU Circ.*, **8656**, 2
- Cochran, A. L., & Cochran, W. D. 1991, *Icarus*, **90**, 172
- Cochran, A. L., Cochran, W. D., & Barker, E. S. 1982, *ApJ*, **254**, 816
- Crovisier, J., Biver, N., Bockelée-Morvan, D., Colom, P., Jorda, L., Lellouch, E., Paubert, G., & Despois, D. 1995, *Icarus*, **115**, 213
- Cruikshank, D. P., et al. 1998, *Icarus*, **135**, 389
- Davies, J. K., Roush, T. L., Cruikshank, D. P., Bartholomew, M. J., Geballe, T. R., Owen, T., & de Bergh, C. 1997, *Icarus*, **127**, 238
- Delsanti, A., Hainaut, O., Jourdeuil, E., Meech, K. J., Boehnhardt, H., & Barrera, L. 2004, *A&A*, **417**, 1145
- Devlin, J. P. 2001, *JGR*, **106**, 33333
- Dotto, E., Barucci, M. A., Leyrat, C., Romon, J., de Bergh, C., & Licandro, J. 2003, *Icarus*, **164**, 122
- Enzian, A., Cabot, H., & Klinger, J. 1997, *A&A*, **319**, 995
- Fernández, Y. R., Jewitt, D. C., & Sheppard, S. S. 2002, *AJ*, **123**, 1050
- Foster, M. J., Green, S. F., McBride, N., & Davies, J. K. 1999, *Icarus*, **141**, 408
- Groussin, O., et al. 2007, *Icarus*, **191**, 63
- Grün, E., et al. 2001, *A&A*, **377**, 1098
- Gunnarsson, M., Bockelée-Morvan, D., Biver, N., Crovisier, J., & Rickman, H. 2008, *A&A*, **484**, 537
- Hahn, G., Lagerkvist, C. I., Karlsson, O., Oja, T., & Stoss, R. M. 2006, *Astron. Nachr.*, **327**, 17
- Hartmann, W. K., Tholen, D. J., Meech, K. J., & Cruikshank, D. P. 1990, *Icarus*, **83**, 1
- Horner, J., Evans, N. W., & Bailey, M. E. 2004, *MNRAS*, **354**, 798
- Jewitt, D. 1990, *ApJ*, **351**, 277
- Jewitt, D., Garland, C. A., & Aussen, H. 2008, *AJ*, **135**, 400
- Jewitt, D., & Luu, J. 1989, *AJ*, **97**, 1766
- Jewitt, D., & Kalas, P. 1998, *ApJ*, **499**, L103

- Jewitt, D. C. 2004, Comets, [II](#), 659
- Jewitt, D. C., & Luu, J. 2004, [Nature](#), **432**, 731
- Kiselev, N. N., & Chernova, G. P. 1979, *SvA Lett.*, **5**, 156
- Klinger, J., Lvasseur-Regourd, A.-C., Bouziani, N., & Enzian, A. 1996, [Planet. Space Sci.](#), **44**, 637
- Kolokolova, L., Hanner, M. S., Lvasseur-Regourd, A.-C., & Gustafson, B. Å. S. 2004, Comets, [II](#), 577
- Kowal, C. T., Liller, W., & Marsden, B. G. 1979, in IAU Symp. 81, Dynamics of the Solar System, ed. R. L. Duncombe (Dordrecht: Kluwer), 245
- Kouchi, A., & Sirono, S.-i. 2001, [Geophys. Res. Lett.](#), **28**, 827
- Kusnirak, P., & Balam, D. 2000, *IAU Circ.*, **7368**, 2
- Lamy, P. L., Toth, I., Fernandez, Y. R., & Weaver, H. A. 2004, Comets, [II](#), 223
- Landolt, A. U. 1992, [AJ](#), **104**, 340
- Luu, J., Jewitt, D., & Cloutis, E. 1994, [Icarus](#), **109**, 133
- Luu, J. X., & Jewitt, D. C. 1990, [AJ](#), **100**, 913
- Luu, J. X., Jewitt, D. C., & Trujillo, C. 2000, [ApJ](#), **531**, L151
- Mazzotta Epifani, E., Palumbo, P., Capria, M. T., Cremonese, G., Fulle, M., & Colangeli, L. 2006, [A&A](#), **460**, 935
- McKinnon, W. B. 2002, in Asteroids, Comets, and Meteors: ACM 2002 (ESA SP-500; Noordwijk: ESA Publications Division), 29
- Meech, K. J., & Belton, M. J. S. 1990, [AJ](#), **100**, 1323
- Meech, K., et al. 2009, *Icarus*, in press
- Merk, R., & Prialnik, D. 2003, [Earth Moon Planets](#), **92**, 359
- Merlin, F., Guilbert, A., Dumas, C., Barucci, M. A., de Bergh, C., & Vernazza, P. 2007, [A&A](#), **466**, 1185
- Morbidelli, A. 2008, Saas-Fee Advanced Course 35: Trans-Neptunian Objects and Comets (Berlin: Springer-Verlag), 79
- Mueller, B. E. A., Tholen, D. J., Hartmann, W. K., & Cruikshank, D. P. 1992, [Icarus](#), **97**, 150
- Notesco, G., & Bar-Nun, A. 1996, [Icarus](#), **122**, 118
- Notesco, G., Bar-Nun, A., & Owen, T. 2003, [Icarus](#), **162**, 183
- Oke, J. B., et al. 1995, [PASP](#), **107**, 375
- Peixinho, N., Doressoundiram, A., Delsanti, A., Boehnhardt, H., Barucci, M. A., & Belskaya, I. 2003, [A&A](#), **410**, L29
- Pinilla Alonso, N., Brunetto, R., Licandro, J., Gil-Hutton, R., Roush, T., & Strazulla, G. 2009, [A&A](#), **496**, 547
- Prialnik, D. 1997, [ApJ](#), **478**, L107
- Prialnik, D., Brosch, N., & Ianovici, D. 1995, *MNRAS*, **276**, 1148
- Roemer, E. 1958, [PASP](#), **70**, 272
- Romon-Martin, J., Delahodde, C., Barucci, M. A., de Bergh, C., & Peixinho, N. 2003, [A&A](#), **400**, 369
- Rousselot, P. 2008, [A&A](#), **480**, 543
- Rousselot, P., Petit, J.-M., Poulet, F., & Sergeev, A. 2005, [Icarus](#), **176**, 478
- Russell, H. N. 1916, [ApJ](#), **43**, 173
- Schmitt, B., Espinasse, S., Grim, R. J. A., Greenberg, J. M., & Klinger, J. 1989, Physics and Mechanics of Cometary Materials (ESA SP-302; Noordwijk: ESA Publications Division), 65
- Senay, M. C., & Jewitt, D. 1994, [Nature](#), **371**, 229
- Stansberry, J. A., et al. 2004, [ApJS](#), **154**, 463
- Szabó, G. M., Kiss, L. L., & Sárneczky, K. 2008, [ApJ](#), **677**, L121
- Tiscareno, M. S., & Malhotra, R. 2003, [AJ](#), **126**, 3122
- Trigo-Rodríguez, J. M., García-Melendo, E., Davidsson, B. J. R., Sánchez, A., Rodríguez, D., Lacruz, J., de Los Reyes, J. A., & Pastor, S. 2008, [A&A](#), **485**, 599
- Trujillo, C. A., Brown, M. E., Barkume, K. M., Schaller, E. L., & Rabinowitz, D. L. 2007, [ApJ](#), **655**, 1172
- Volk, K., & Malhotra, R. 2008, [ApJ](#), **687**, 714
- Weissman, P. R., Chesley, S. R., Choi, Y. J., Bauer, J. M., Tegler, S. C., Romanishin, W. J., Consolmagno, G., & Stansberry, J. A. 2006, *BAAS*, **38**, 551
- Whipple, F. L. 1950, [ApJ](#), **111**, 375
- Womack, M., & Stern, S. A. 1999, *Astron. Vestn.*, **33**, 216